Outline: A brief history of Nuclear Astrophysics N. Prantzos

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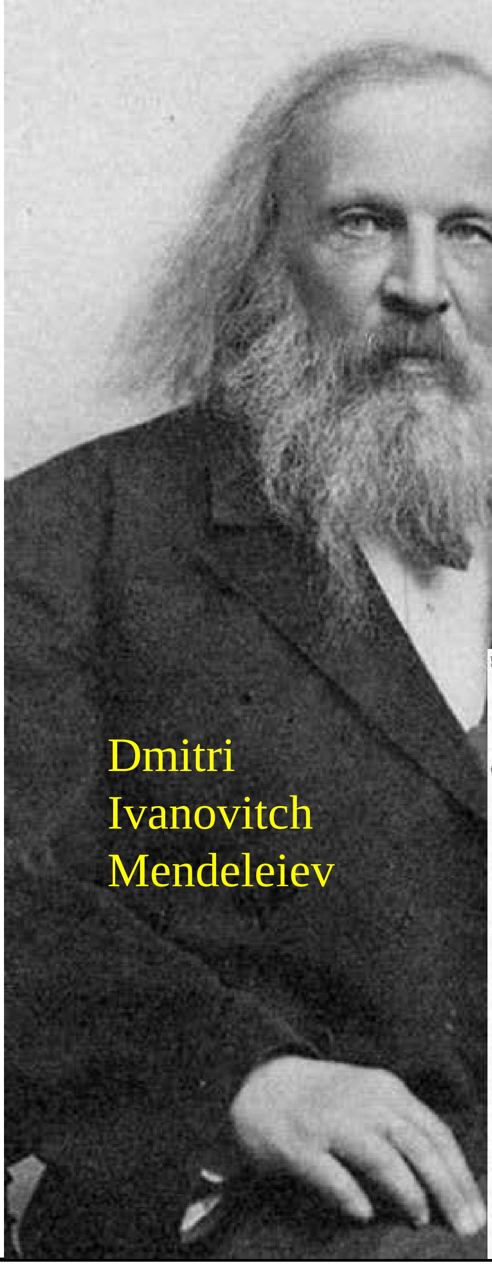
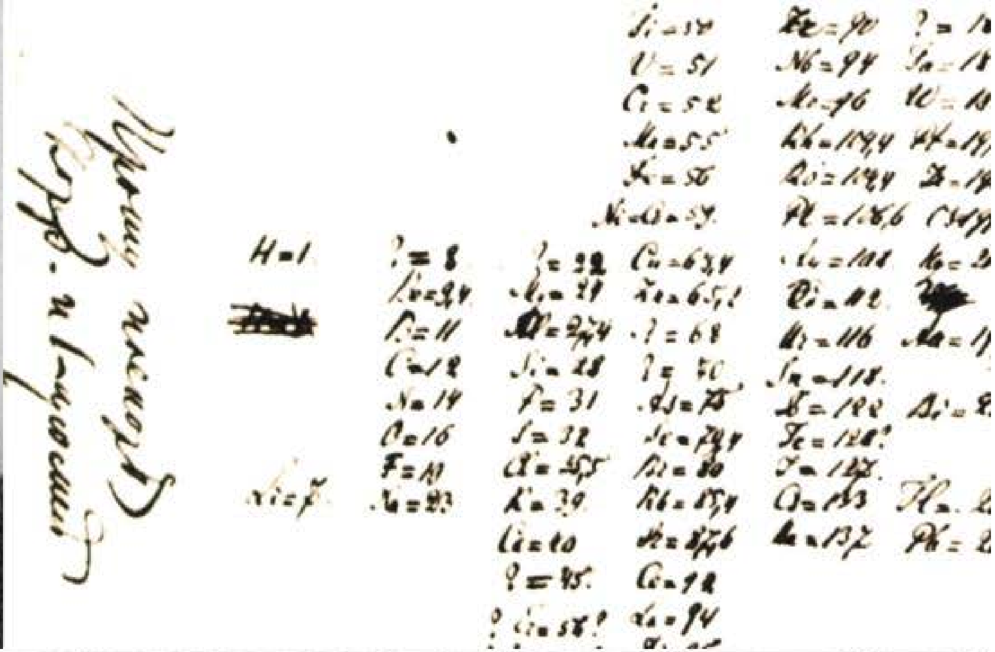
Kernphysik Skript (H.J. Wollersheim)

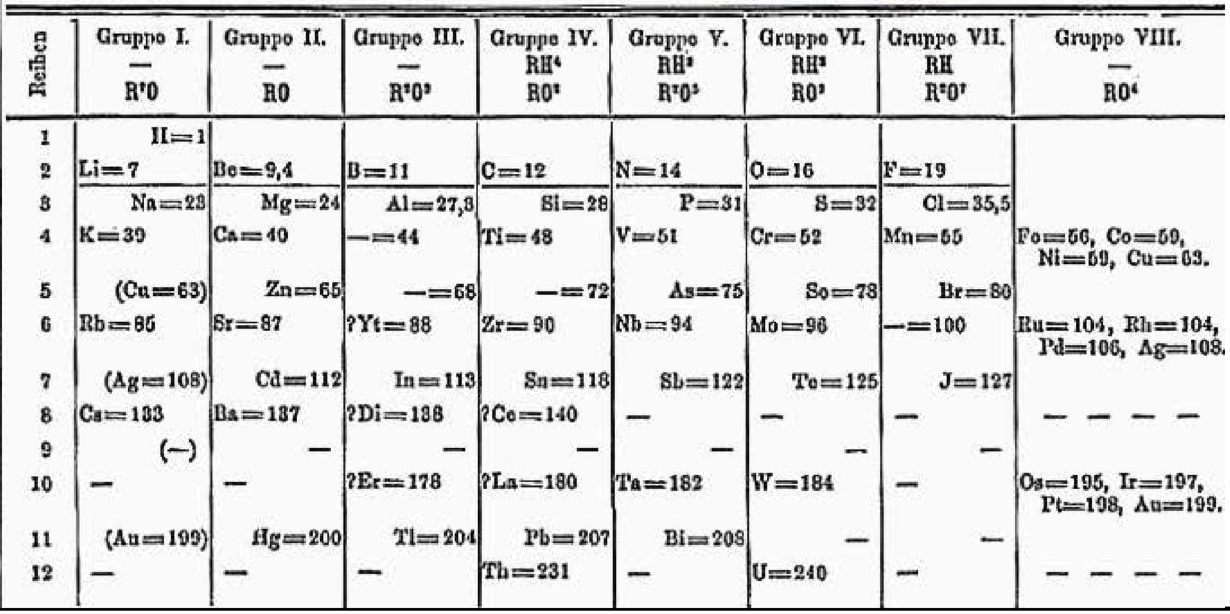
1. solar abundances of the elements
2. nuclear reactions with Hydrogen
3. expanding Universe
4. formation of ^{12}C
5. explosive nucleosynthesis (supernovae)

The origin of the elements table of the elements



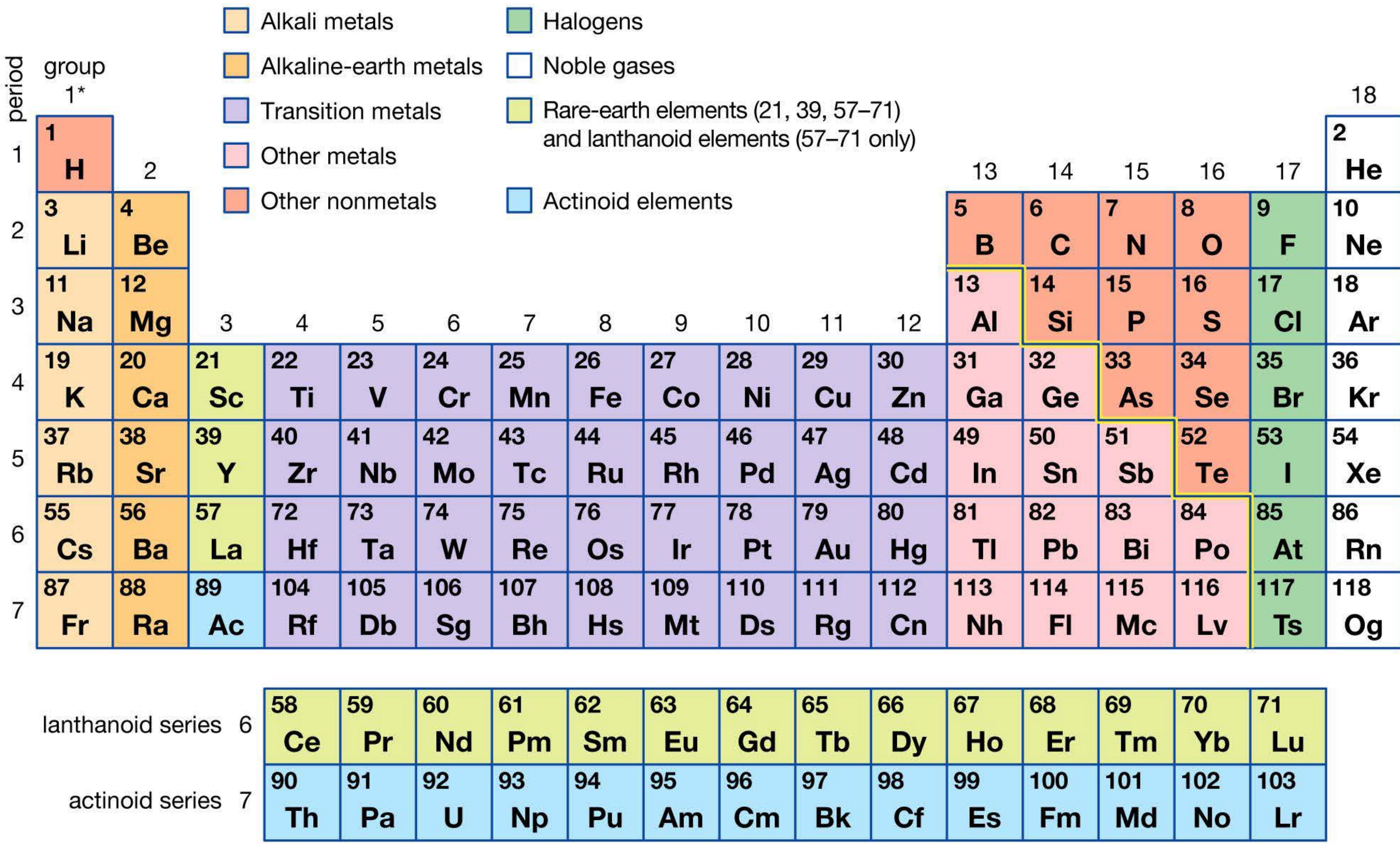
Dmitri
Ivanovitch
Mendeleiev





Reihen	Gruppo I. — R^2O	Gruppo II. — RO	Gruppo III. — R^2O^3	Gruppo IV. RH^4 RO^2	Gruppo V. RH^3 R^2O^5	Gruppo VI. RH^2 RO^3	Gruppo VII. RH R^2O^7	Gruppo VIII. — RO^4
1	H=1							
2	Li=7	Be=9,4	B=11	C=12	N=14	O=16	F=19	
3	Na=23	Mg=24	Al=27,3	Si=28	P=31	S=32	Cl=35,5	
4	K=39	Ca=40	—=44	Ti=48	V=51	Cr=52	Mn=55	Fe=56, Co=59, Ni=59, Cu=63.
5	(Cu=63)	Zn=65	—=68	—=72	As=75	Se=78	Br=80	
6	Rb=85	Sr=87	?Yt=88	Zr=90	Nb=94	Mo=96	—=100	Ru=104, Rh=104, Pd=106, Ag=108.
7	(Ag=108)	Cd=112	In=113	Sn=118	Sb=122	Te=125	J=127	
8	Cs=133	Ba=137	?Di=138	?Ce=140	—	—	—	— — — —
9	(—)	—	—	—	—	—	—	
10	—	—	?Er=178	?La=180	Ta=182	W=184	—	Os=195, Ir=197, Pt=198, Au=199.
11	(Au=199)	Hg=200	Tl=204	Pb=207	Bi=208	—	—	
12	—	—	—	Th=231	—	U=240	—	— — — —

Periodic table of the elements



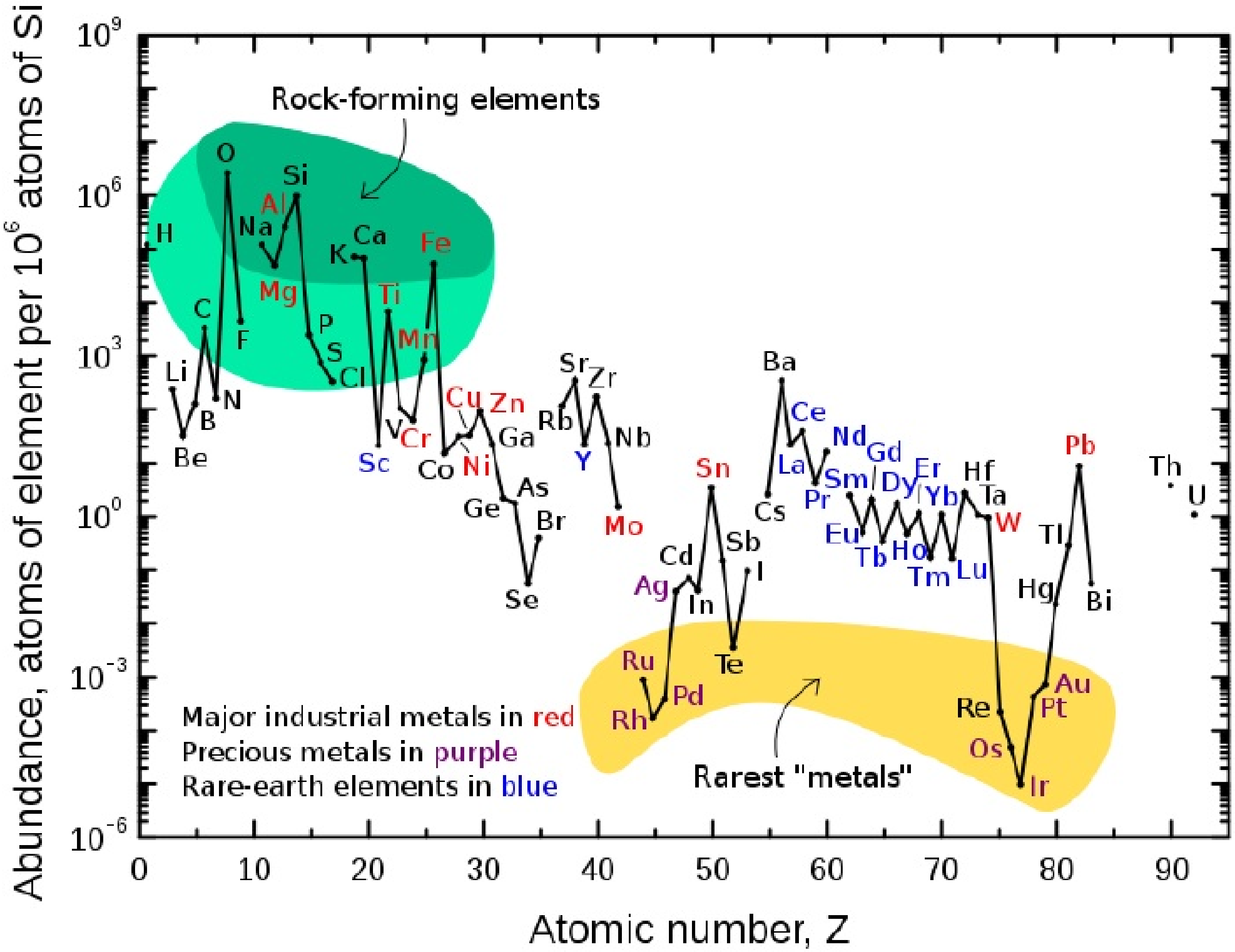
- ☐ Alkali metals
- ☐ Alkaline-earth metals
- ☐ Transition metals
- ☐ Other metals
- ☐ Other nonmetals
- ☐ Halogens
- ☐ Noble gases
- ☐ Rare-earth elements (21, 39, 57–71) and lanthanoid elements (57–71 only)
- ☐ Actinoid elements

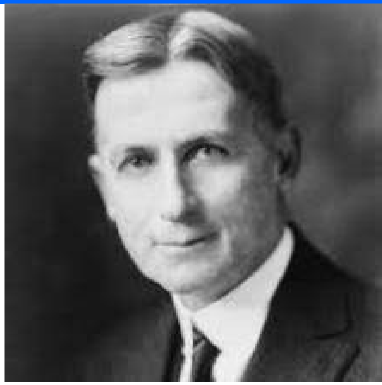
period	group 1*	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18
1	1 **H**																	2 **He**
2	3 **Li**	4 **Be**											5 **B**	6 **C**	7 **N**	8 **O**	9 **F**	10 **Ne**
3	11 **Na**	12 **Mg**											13 **Al**	14 **Si**	15 **P**	16 **S**	17 **Cl**	18 **Ar**
4	19 **K**	20 **Ca**	21 **Sc**	22 **Ti**	23 **V**	24 **Cr**	25 **Mn**	26 **Fe**	27 **Co**	28 **Ni**	29 **Cu**	30 **Zn**	31 **Ga**	32 **Ge**	33 **As**	34 **Se**	35 **Br**	36 **Kr**
5	37 **Rb**	38 **Sr**	39 **Y**	40 **Zr**	41 **Nb**	42 **Mo**	43 **Tc**	44 **Ru**	45 **Rh**	46 **Pd**	47 **Ag**	48 **Cd**	49 **In**	50 **Sn**	51 **Sb**	52 **Te**	53 **I**	54 **Xe**
6	55 **Cs**	56 **Ba**	57 **La**	72 **Hf**	73 **Ta**	74 **W**	75 **Re**	76 **Os**	77 **Ir**	78 **Pt**	79 **Au**	80 **Hg**	81 **Tl**	82 **Pb**	83 **Bi**	84 **Po**	85 **At**	86 **Rn**
7	87 **Fr**	88 **Ra**	89 **Ac**	104 **Rf**	105 **Db**	106 **Sg**	107 **Bh**	108 **Hs**	109 **Mt**	110 **Ds**	111 **Rg**	112 **Cn**	113 **Nh**	114 **Fl**	115 **Mc**	116 **Lv**	117 **Ts**	118 **Og**

lanthanoid series 6	58 **Ce**	59 **Pr**	60 **Nd**	61 **Pm**	62 **Sm**	63 **Eu**	64 **Gd**	65 **Tb**	66 **Dy**	67 **Ho**	68 **Er**	69 **Tm**	70 **Yb**	71 **Lu**
actinoid series 7	90 **Th**	91 **Pa**	92 **U**	93 **Np**	94 **Pu**	95 **Am**	96 **Cm**	97 **Bk**	98 **Cf**	99 **Es**	100 **Fm**	101 **Md**	102 **No**	103 **Lr**

*Numbering system adopted by the International Union of Pure and Applied Chemistry (IUPAC).

Composition of Earth´s crust



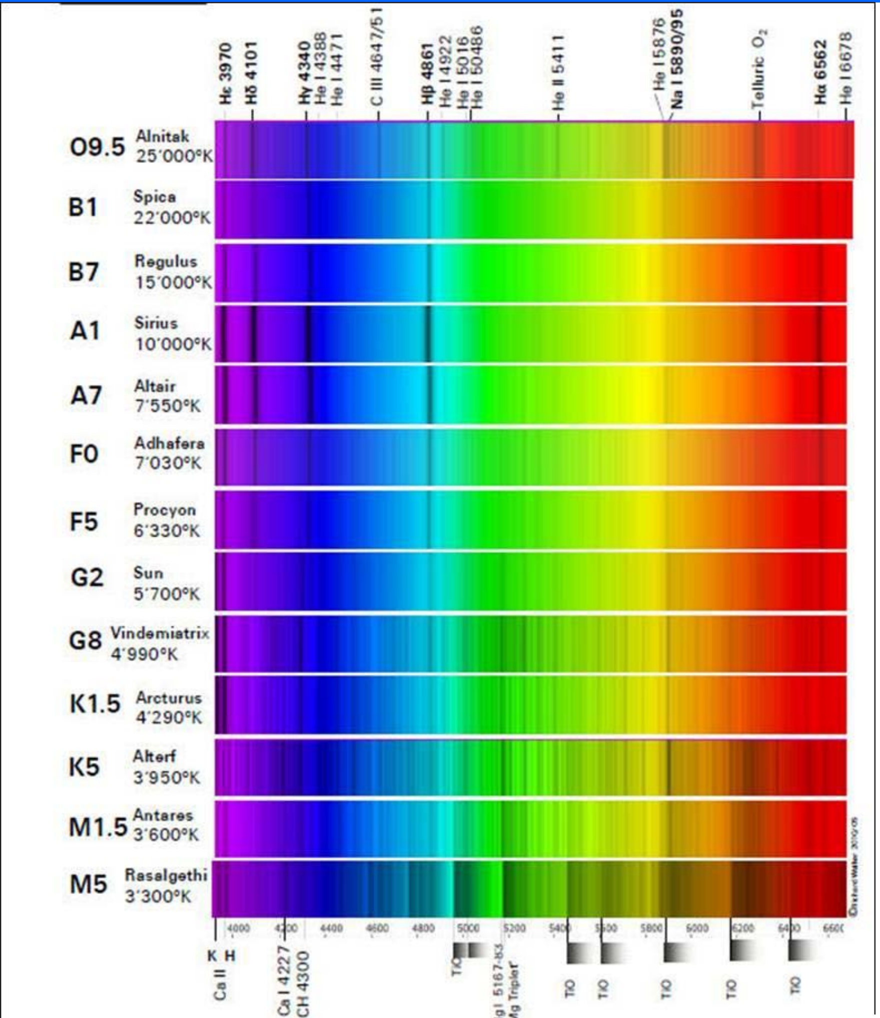


William Draper Harkins

Harkins rule (1915)
elements with specific properties are more abundant than others:
even vs odd charge

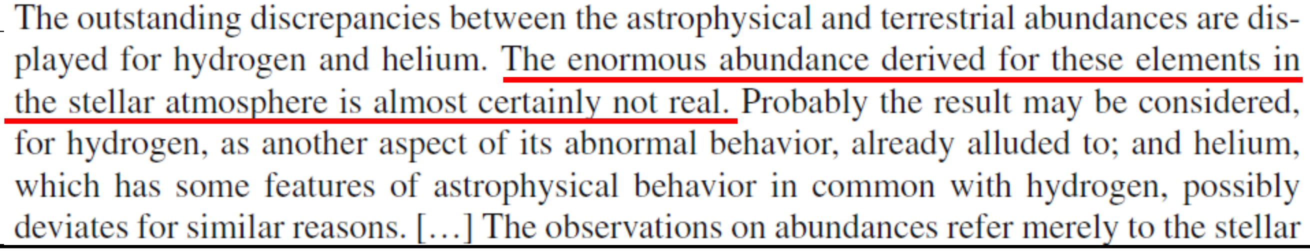
It should be related to the **structure of their nuclei**, and not to their position in the Periodic Table

Stellar spectroscopy reveals the presence of chemical elements in stellar surfaces



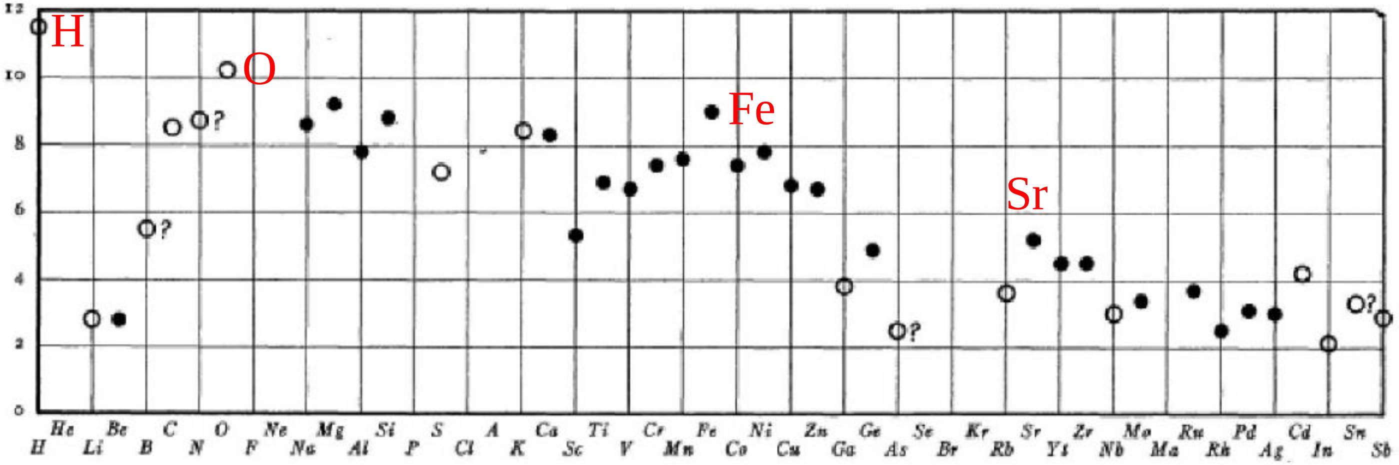
Cecilia Payne

1925: H and He are the most abundant elements in stellar atmospheres



The outstanding discrepancies between the astrophysical and terrestrial abundances are displayed for hydrogen and helium. The enormous abundance derived for these elements in the stellar atmosphere is almost certainly not real. Probably the result may be considered, for hydrogen, as another aspect of its abnormal behavior, already alluded to; and helium, which has some features of astrophysical behavior in common with hydrogen, possibly deviates for similar reasons. […] The observations on abundances refer merely to the stellar

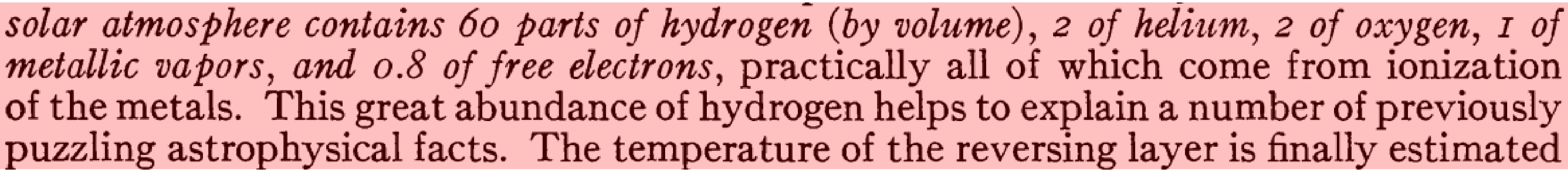
On the composition of the Sun´s atmosphere



Henry Norris Russell

1929:

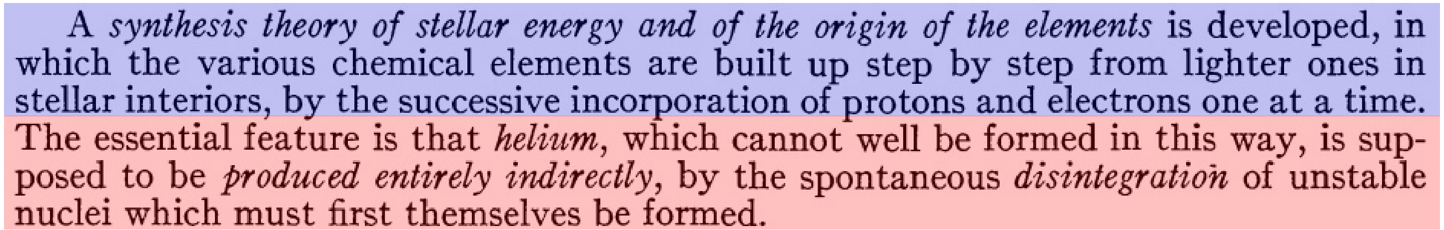
Harkins rule (1915)
elements with specific properties are more abundant than others:
even vs odd charge



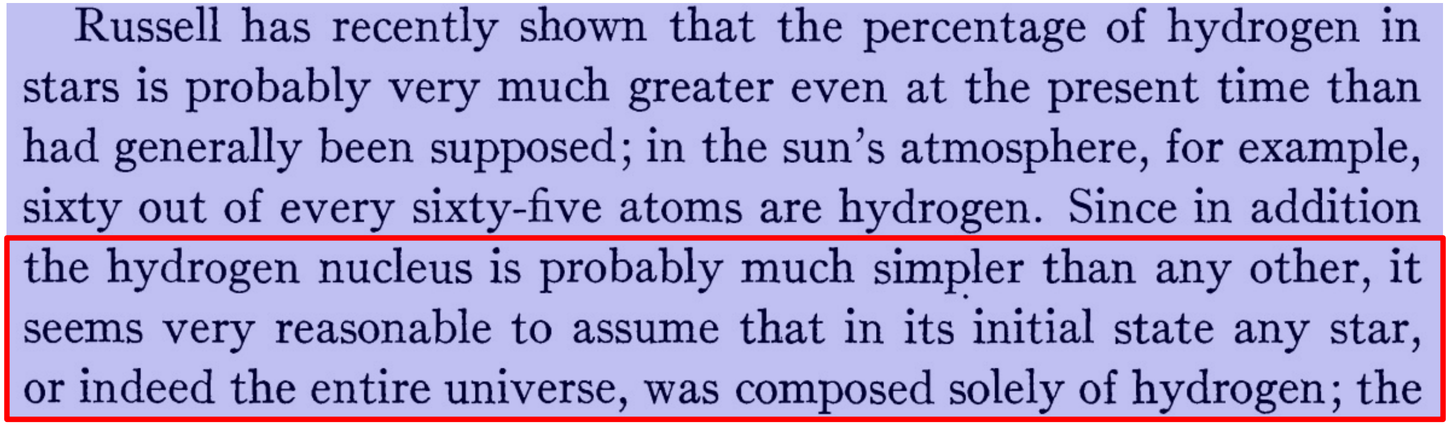
solar atmosphere contains 60 parts of hydrogen (by volume), 2 of helium, 2 of oxygen, 1 of metallic vapors, and 0.8 of free electrons, practically all of which come from ionization of the metals. This great abundance of hydrogen helps to explain a number of previously puzzling astrophysical facts. The temperature of the reversing layer is finally estimated

Atomic synthesis and stellar energy

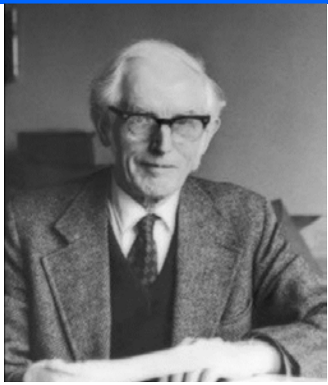
1931:



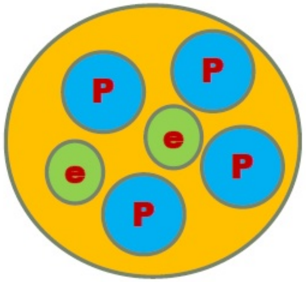
A *synthesis theory of stellar energy and of the origin of the elements* is developed, in which the various chemical elements are built up step by step from lighter ones in stellar interiors, by the successive incorporation of protons and electrons one at a time. The essential feature is that *helium*, which cannot well be formed in this way, is supposed to be *produced entirely indirectly*, by the spontaneous *disintegration* of unstable nuclei which must first themselves be formed.



Russell has recently shown that the percentage of hydrogen in stars is probably very much greater even at the present time than had generally been supposed; in the sun's atmosphere, for example, sixty out of every sixty-five atoms are hydrogen. Since in addition the hydrogen nucleus is probably much simpler than any other, it seems very reasonable to assume that in its initial state any star, or indeed the entire universe, was composed solely of hydrogen; the



Robert D. Atkinson



$$4p^{+} + 2e^{-} \Longrightarrow He$$

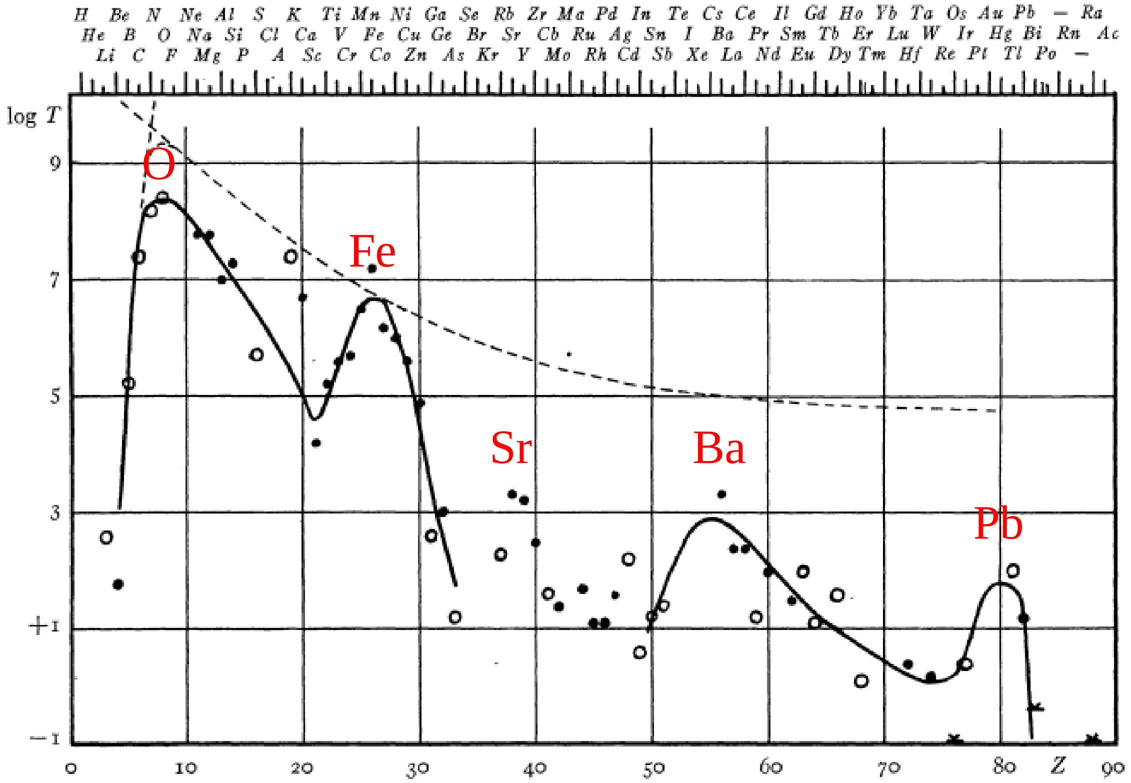
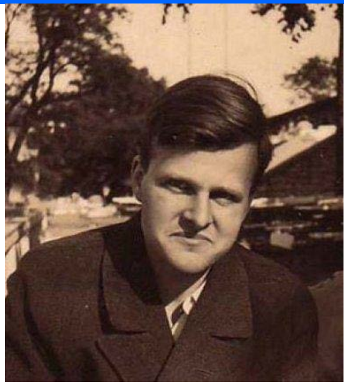


FIG. 2.—Amount of the elements in the sun's atmosphere. (After Russell; ordinates at odd Z values increased by 0.6.) - - - - Equilibrium amounts; • First class determinations; ○ Second class determinations.

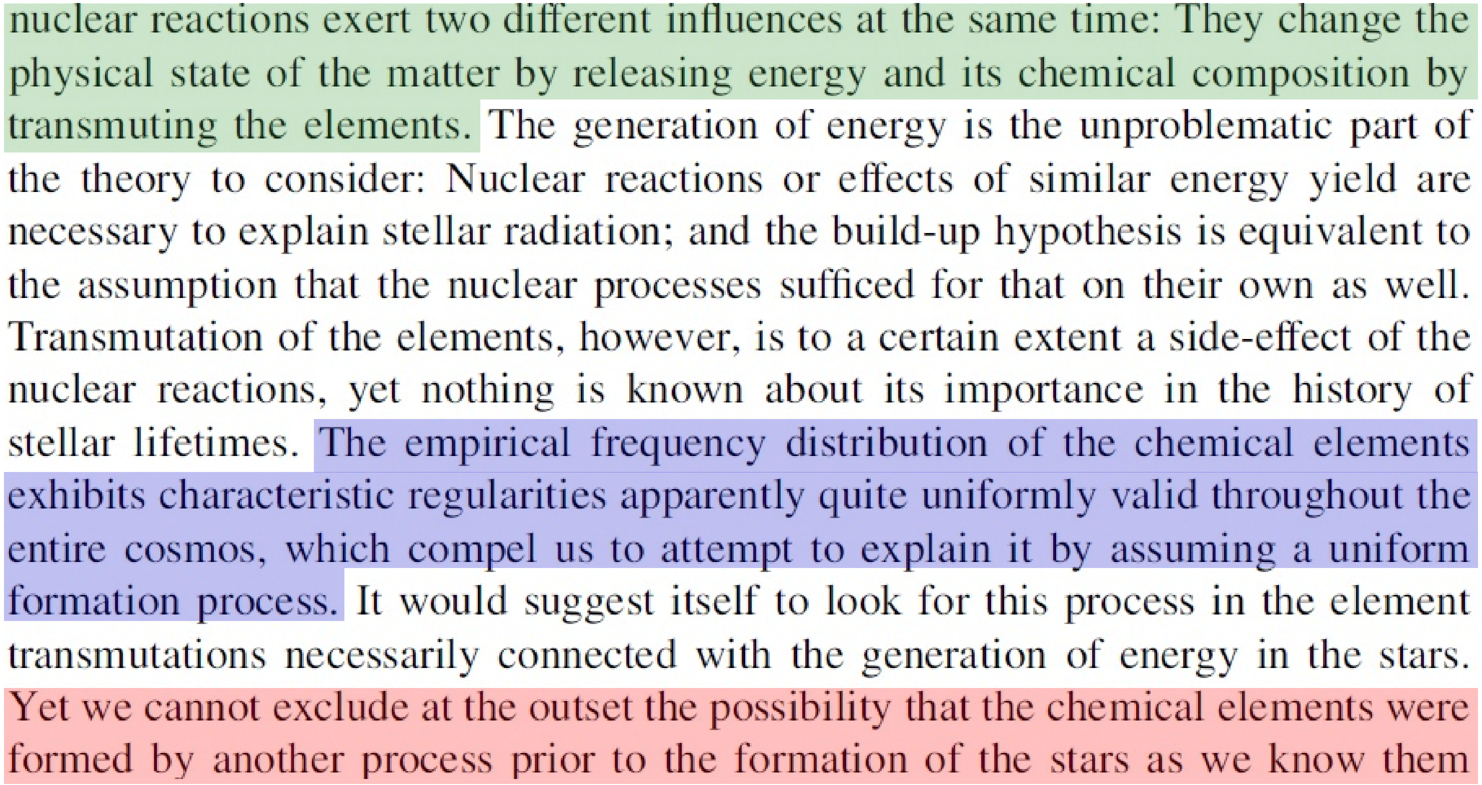
The *relative proportions of the elements* in stars of the main sequence follow from the theory, in *excellent qualitative agreement* with Russell's figures for the sun. The scarcity of the lightest elements, the principal maximum at a fairly early point, a minimum before the iron group, a maximum in it, a scarcity of all elements above it, and minor maxima in the barium and lead regions all follow (Fig. 2) without any special assumptions, from Gamow's theory of nuclear stability, owing to the peculiarities of the Aston mass-defect curve.

On elementary transmutations in the interior of stars

Are stars making their own elements or is it something else preceding them?



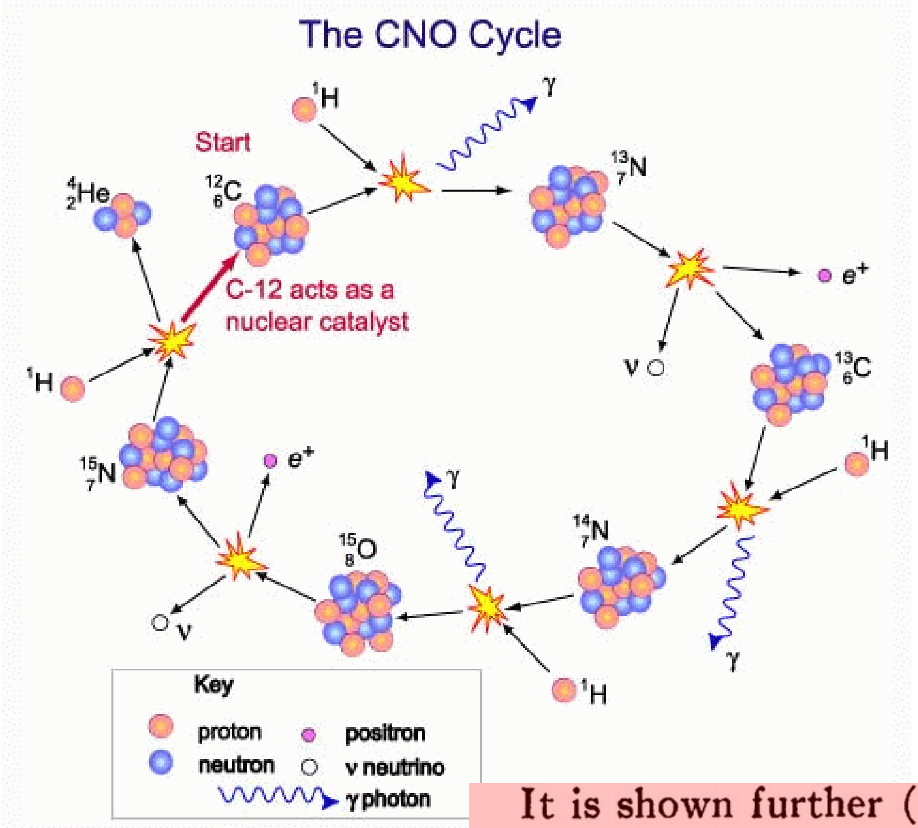
Carl Friedrich
von Weizäcker



nuclear reactions exert two different influences at the same time: They change the physical state of the matter by releasing energy and its chemical composition by transmuting the elements. The generation of energy is the unproblematic part of the theory to consider: Nuclear reactions or effects of similar energy yield are necessary to explain stellar radiation; and the build-up hypothesis is equivalent to the assumption that the nuclear processes sufficed for that on their own as well. Transmutation of the elements, however, is to a certain extent a side-effect of the nuclear reactions, yet nothing is known about its importance in the history of stellar lifetimes. The empirical frequency distribution of the chemical elements exhibits characteristic regularities apparently quite uniformly valid throughout the entire cosmos, which compel us to attempt to explain it by assuming a uniform formation process. It would suggest itself to look for this process in the element transmutations necessarily connected with the generation of energy in the stars. Yet we cannot exclude at the outset the possibility that the chemical elements were formed by another process prior to the formation of the stars as we know them

Energy production in stars

H. A. Bethe



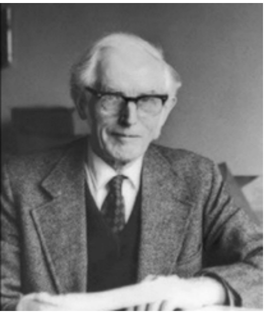
It is shown further (§5–6) that *no elements heavier than* He^4 *can be built up in ordinary stars*. This is due to the fact, mentioned above, that all elements up to boron are disintegrated by proton bombardment (α-emission!) rather than built up (by radiative capture). The instability of Be^8 reduces the formation of heavier elements still further. The production of neutrons in stars is likewise negligible. The heavier elements found in stars must therefore have existed already when the star was formed.

Early studies and ideas about the origin of the elements up to 1940

1925: Cecilia Payne
stars are made mostly of Hydrogen

1915: William Harkins
Even elements on Earth´s crust are more abundant than odd ones to the structure of their nuclei

W. D. Harkins



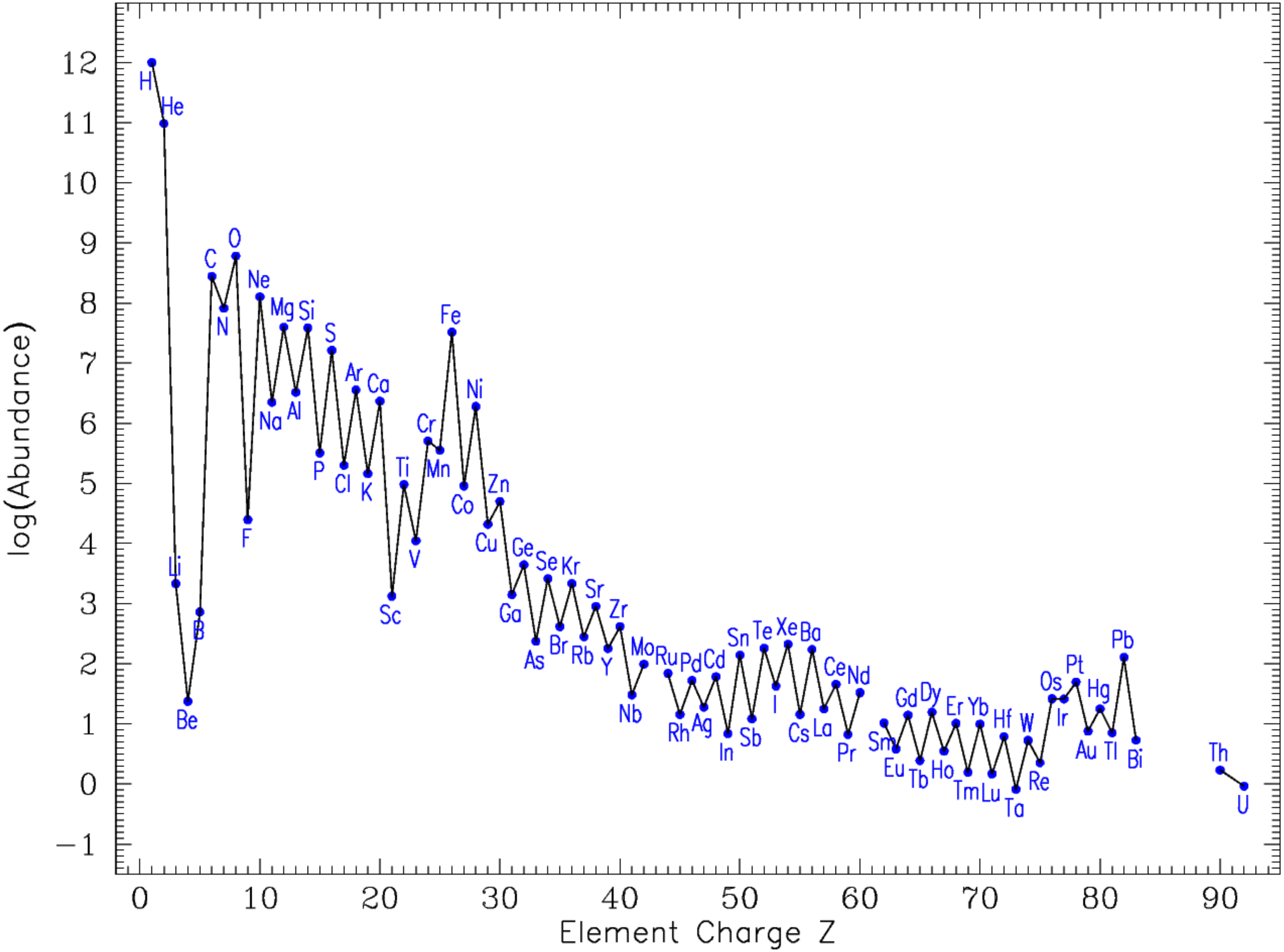
1931: Robert Atkinson
stars build up their elements in their interiors by pre-existing Hydrogen (except for He)

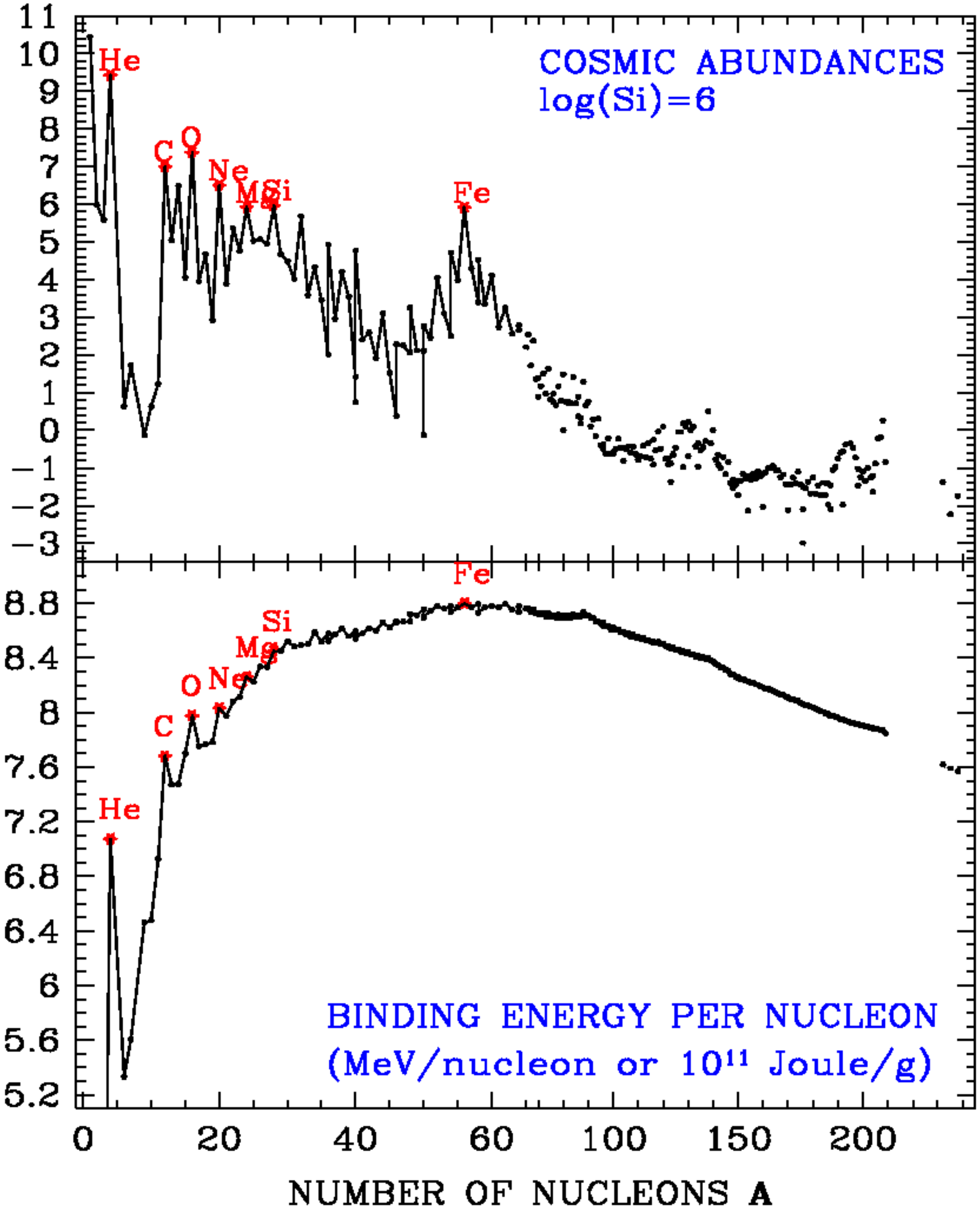
1929: Henri Norris Russell
stellar abundances display specific regularities similar to those on Earth

1939: Hans Bethe
normal stars cannot built in their interiors elements heavier than Helium

1937: Carl F. von Weizsäcker
may be; but perhaps the elements were made by another process, prior to their formation

Solar (Cosmic) Abundances of the elements





Cosmic abundances of nuclides are locally correlated with

nuclear stability

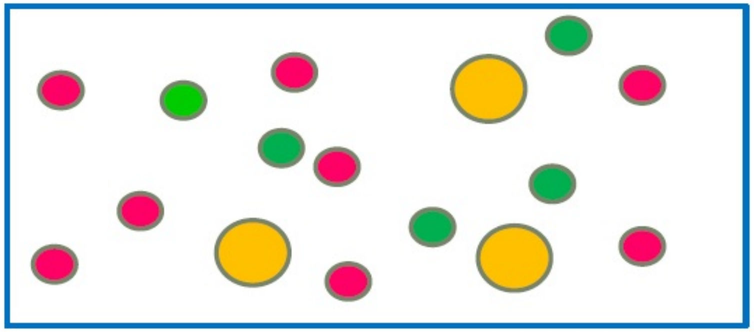
(binding energy per nucleon)

α-nuclei (A = mult of 4), „magic“ nuclei,
Fe peak nuclei
or nuclei with even A or Z
are more abundant than their neighbors

Nuclear processes have shaped the cosmic abundances of the chemical elements

where? **how?**

Nuclear reactions



System of interacting nuclei at temperature T and density ρ evolving for lapse of time t

Nuclear reactions proceeding at rate R $A + B \Longrightarrow C + D$

1. $R \gg 1/t$: All direct and inverse reactions proceed very fast: $A + B \Leftrightarrow C + D$

nuclear statistical equilibrium

$$X_i(A_i, Z_i, T, \rho) = \frac{A}{N_A \cdot \rho} \cdot \omega(T) \cdot \left(\frac{2\pi \cdot k \cdot T \cdot M(A_i, Z_i)}{h^2}\right)^{3/2} \cdot exp\left[\frac{\mu(A_i, Z_i) + \boxed{B(A_i, Z_i)}}{k \cdot T}\right]$$

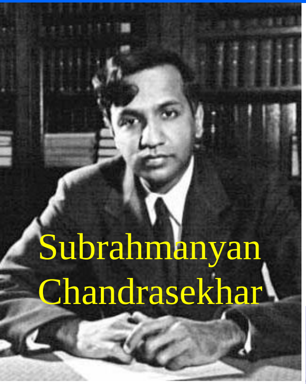
Abundances of nuclei depend on T, ρ and their binding energies $B(A_i, Z_i)$

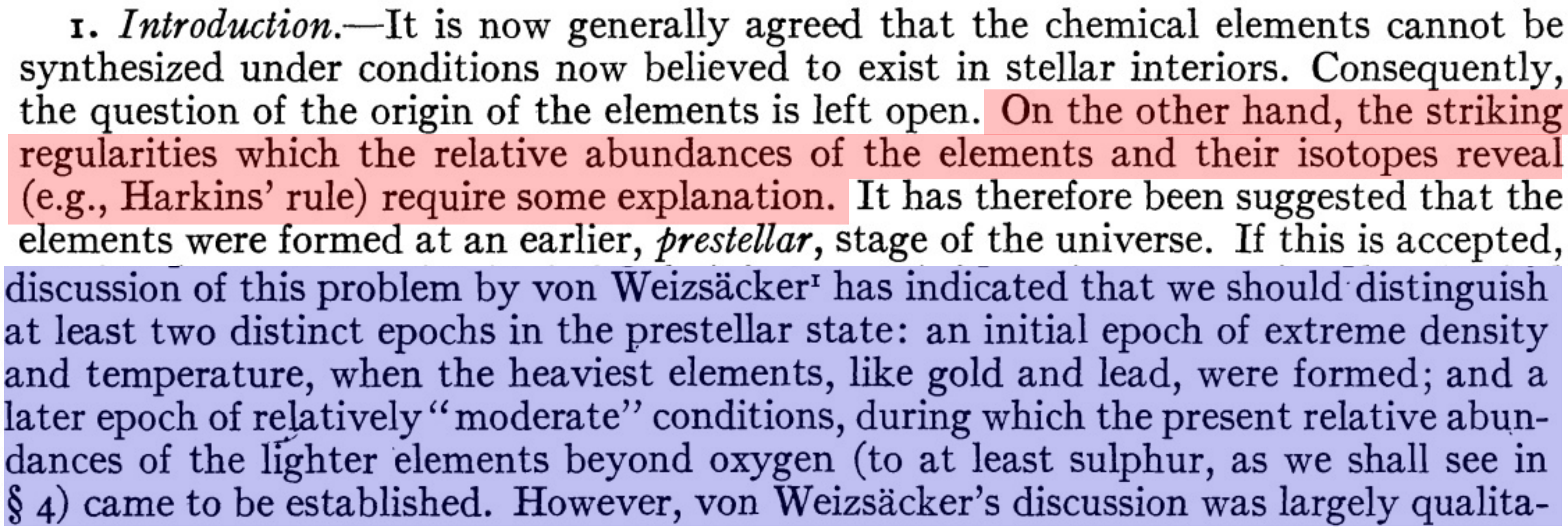
2. $R \leq 1/t$: reactions proceed slowly: $A + B \Rightarrow C + D$

Abundances of nuclei depend on reaction rate $R(T, \rho)$ and are coupled to the abundances of all other nuclei of the system

time-dependent treatment required

An attempt to interpret the relative abundances of the elements 1942


Subrahmanyan Chandrasekhar

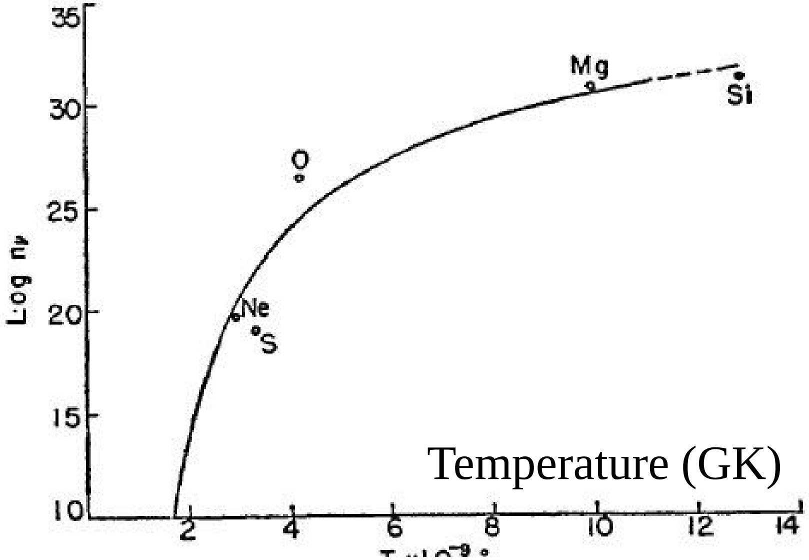

1. *Introduction.*—It is now generally agreed that the chemical elements cannot be synthesized under conditions now believed to exist in stellar interiors. Consequently, the question of the origin of the elements is left open. On the other hand, the striking regularities which the relative abundances of the elements and their isotopes reveal (e.g., Harkins' rule) require some explanation. It has therefore been suggested that the elements were formed at an earlier, *prestellar*, stage of the universe. If this is accepted, discussion of this problem by von Weizsäcker[1] has indicated that we should distinguish at least two distinct epochs in the prestellar state: an initial epoch of extreme density and temperature, when the heaviest elements, like gold and lead, were formed; and a later epoch of relatively "moderate" conditions, during which the present relative abundances of the lighter elements beyond oxygen (to at least sulphur, as we shall see in § 4) came to be established. However, von Weizsäcker's discussion was largely qualita-

Starting at temperature T~10 GK and density ρ~10^8 g/cm^3 built nuclei around Si in conditions of nuclear equilibrium

$$A + B \Leftrightarrow C + D$$

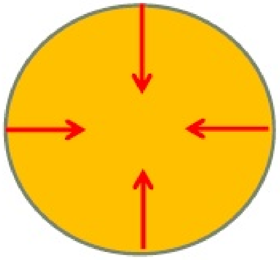
then at lower T and ρ built lighter nuclei

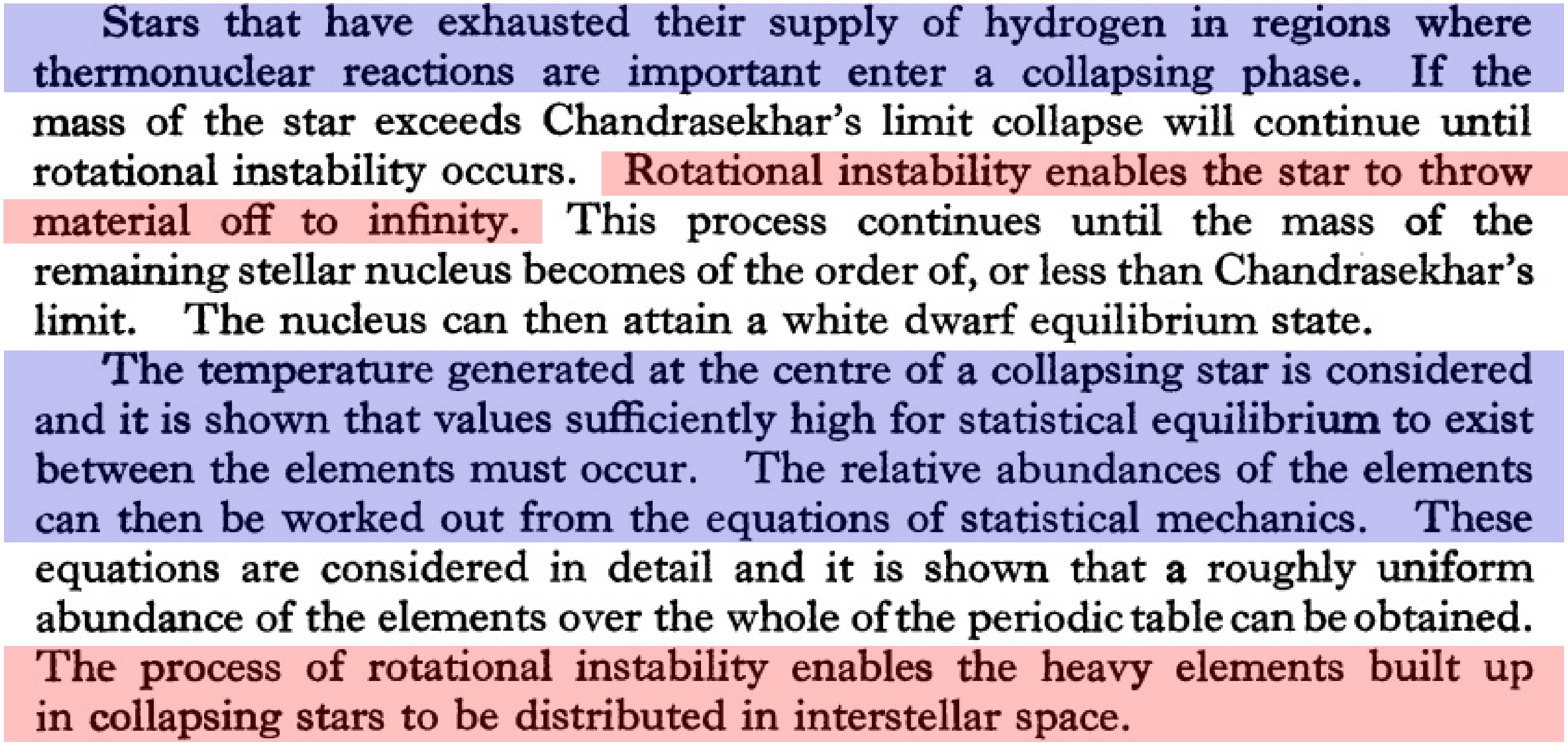
But **Fe** and heavier nuclei **never** produced



The synthesis of the elements from Hydrogen 1946

F. Hoyle



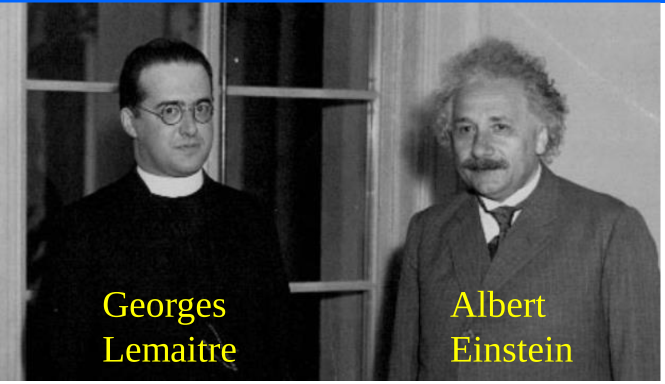


Stars that have exhausted their supply of hydrogen in regions where thermonuclear reactions are important enter a collapsing phase. If the mass of the star exceeds Chandrasekhar's limit collapse will continue until rotational instability occurs. Rotational instability enables the star to throw material off to infinity. This process continues until the mass of the remaining stellar nucleus becomes of the order of, or less than Chandrasekhar's limit. The nucleus can then attain a white dwarf equilibrium state.

The temperature generated at the centre of a collapsing star is considered and it is shown that values sufficiently high for statistical equilibrium to exist between the elements must occur. The relative abundances of the elements can then be worked out from the equations of statistical mechanics. These equations are considered in detail and it is shown that a roughly uniform abundance of the elements over the whole of the periodic table can be obtained. The process of rotational instability enables the heavy elements built up in collapsing stars to be distributed in interstellar space.

But HOW to get material out of the stars?

Expanding Universe and the origin of elements G. Gamow 1946



Lemaitre 1927:

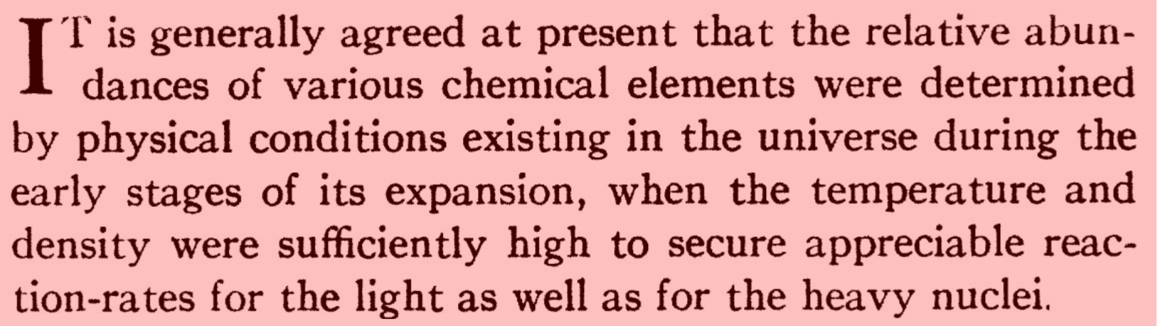
Recession of galaxies explained as due to expansion of the Universe

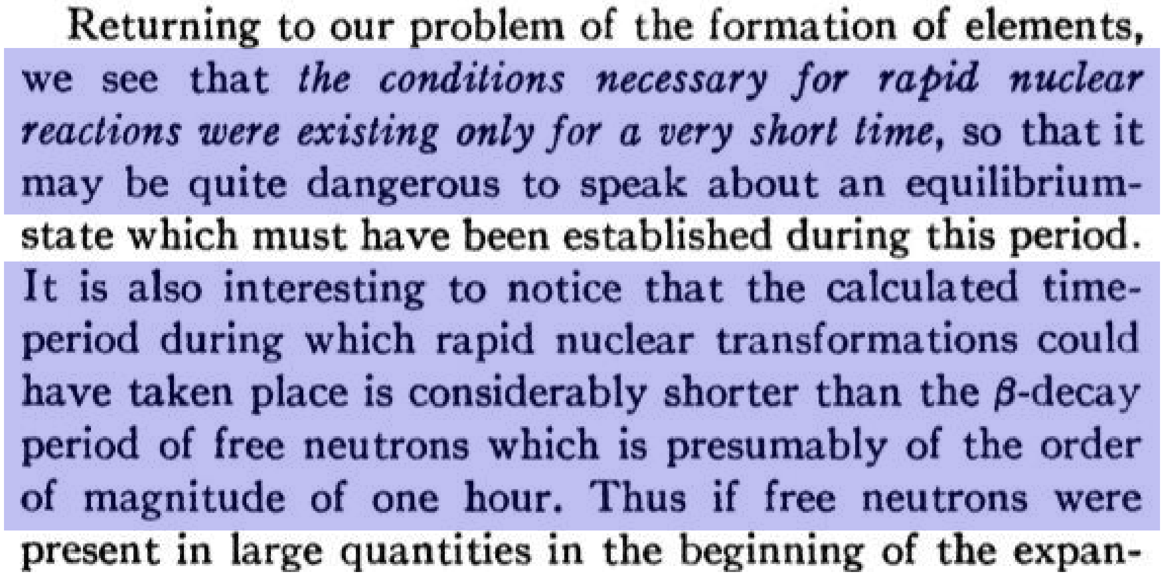
Einstein to Lemaitre:

your calculations are correct, but your physics is atrocious

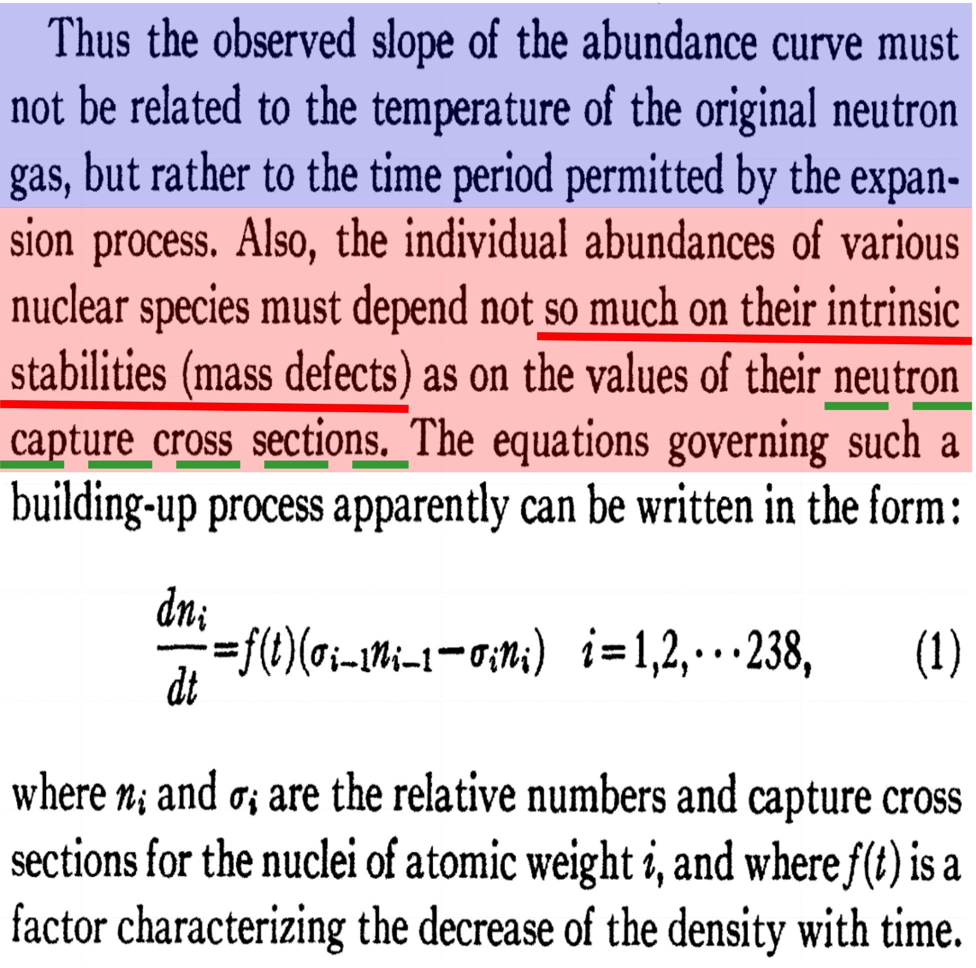
1931:

Explosion of the primeval atom


IT is generally agreed at present that the relative abundances of various chemical elements were determined by physical conditions existing in the universe during the early stages of its expansion, when the temperature and density were sufficiently high to secure appreciable reaction-rates for the light as well as for the heavy nuclei.


Returning to our problem of the formation of elements, we see that *the conditions necessary for rapid nuclear reactions were existing only for a very short time*, so that it may be quite dangerous to speak about an equilibrium-state which must have been established during this period. It is also interesting to notice that the calculated time-period during which rapid nuclear transformations could have taken place is considerably shorter than the β-decay period of free neutrons which is presumably of the order of magnitude of one hour. Thus if free neutrons were present in large quantities in the beginning of the expan-

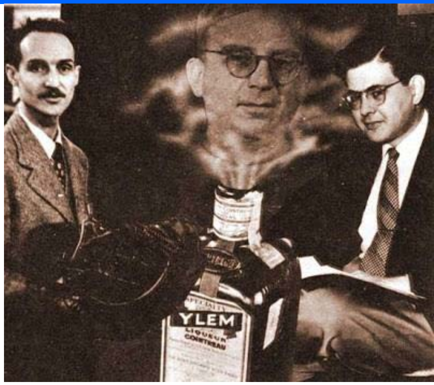
The origin of chemical elements R. A. Alpher, H. Bethe, G. Gamow 1948



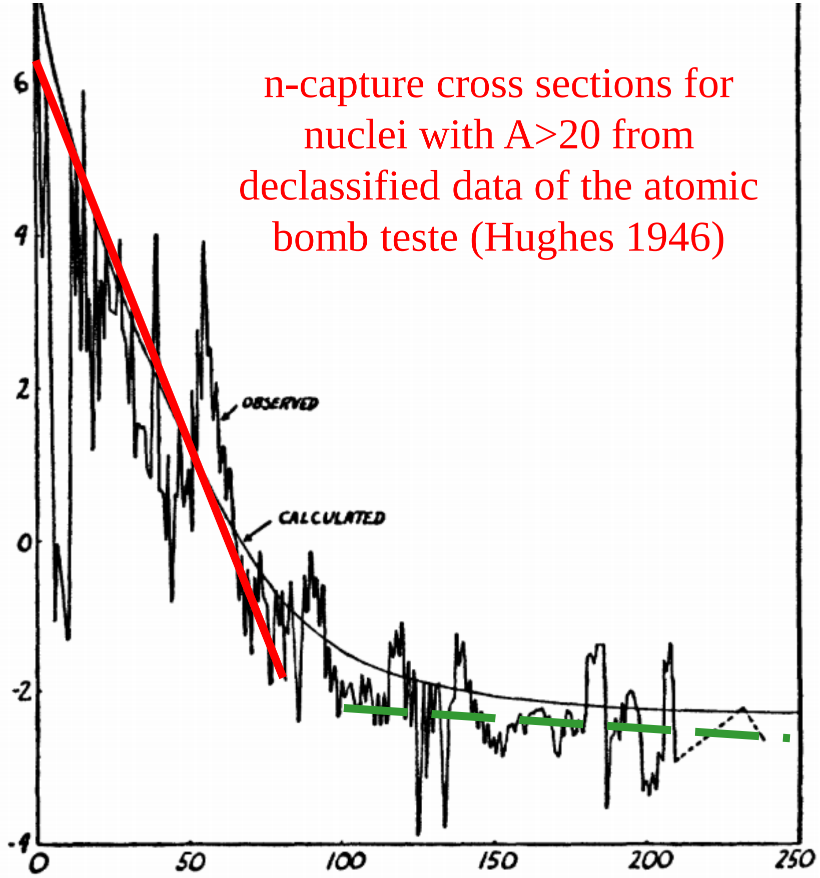
Thus the observed slope of the abundance curve must not be related to the temperature of the original neutron gas, but rather to the time period permitted by the expansion process. Also, the individual abundances of various nuclear species must depend not so much on their intrinsic stabilities (mass defects) as on the values of their neutron capture cross sections. The equations governing such a building-up process apparently can be written in the form:

$$\frac{dn_i}{dt}=f(t)(\sigma_{i-1}n_{i-1}-\sigma_i n_i)\quad i=1,2,\cdots 238, \qquad (1)$$

where n_i and σ_i are the relative numbers and capture cross sections for the nuclei of atomic weight i, and where $f(t)$ is a factor characterizing the decrease of the density with time.



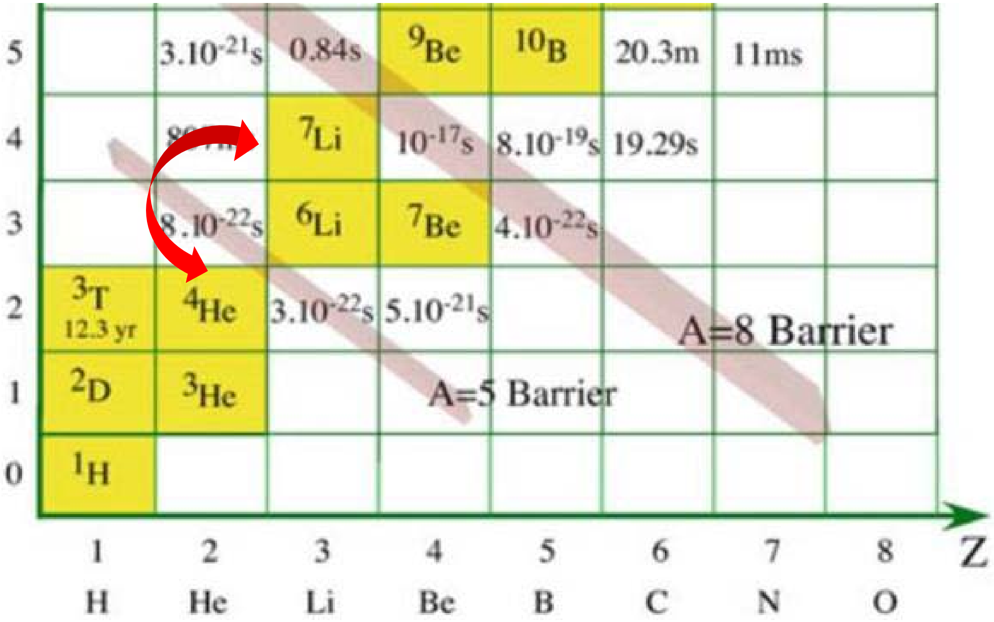
log (abundances)



n-capture cross sections for nuclei with A>20 from declassified data of the atomic bomb teste (Hughes 1946)

Washington Post 16.4.1948: World began in 5 minutes, new theory

At the beginning of everything, the Universe had infinite density concentrated in a single point. Then just 300 seconds – 5 minutes – after the start of everything, there was a rapid expansion and cooling of the primordial matter. The neutrons – those are the particles that trigger the atomic bomb – started decaying into protons and building up heavier chemical elements … This act of creation of the chemical elements took the surprising short time of an hour. (the Bible story said something about 6 days for the act of creation)



Fermi and Turkevich (1949)
no elements beyond He, (Li) because of A=5 (8) gap

Hayashi (1950)
at T~10^{10} K: $n \Leftrightarrow p$ *equilibrium*

after the discovery of 3 K microwave background (Penzias and Wilson 1965)

Peebles 1966, Wagoner, Fowler, Hoyle 1967
first calculation of realistic Big Bang nucleosyntheses

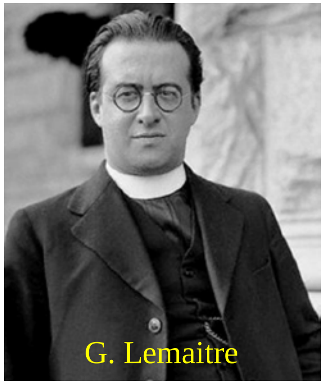
Address to the Pontifical Academy of Sciences 1951

The proof for the existence of God in the light of modern natural science

Present day science, with one sweeping step back across millions of centuries, has succeeded in bearing witness to that primordial “Fiat lux” uttered at the moment when, along with matter, there burst forth from nothing a sea of light and radiation, while the particles of chemical elements split and formed into millions of galaxies …

Hence, creation took place in time. Therefore, there is a Creator. Therefore,, God exists! Although it is neither explicit nor complete, this is the reply we were awaiting from science, and which the present human generation is awaiting from it

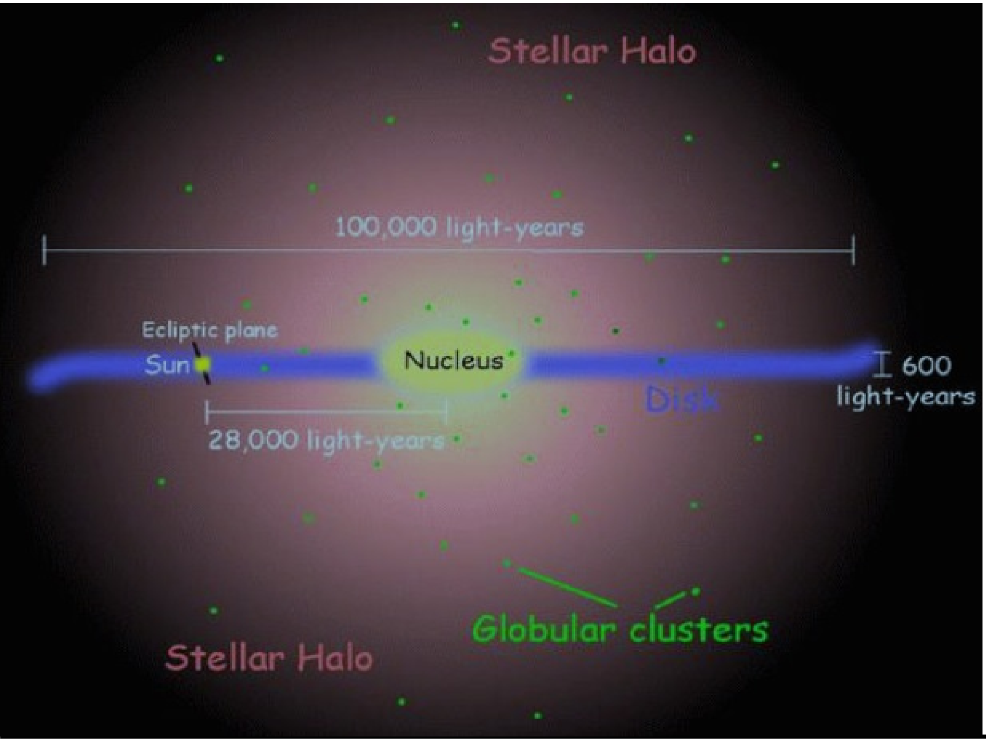
We may speak of this event as of a beginning. I do not say a creation. Physically it is a beginning in the sense that if something happened before, it has no observable influence on the behavior of our universe, as any feature of matter before this beginning has been completely lost by the extreme contraction at the theoretical zero. The question if it was really a beginning or rather a creation, something started from nothing, is a philosophical question which cannot be settled by physical or astronomical considerations



G. Gamow:
all elements were produced in the hot primordial Universe (Big Bang) by successive neutron captures

Early 1950s

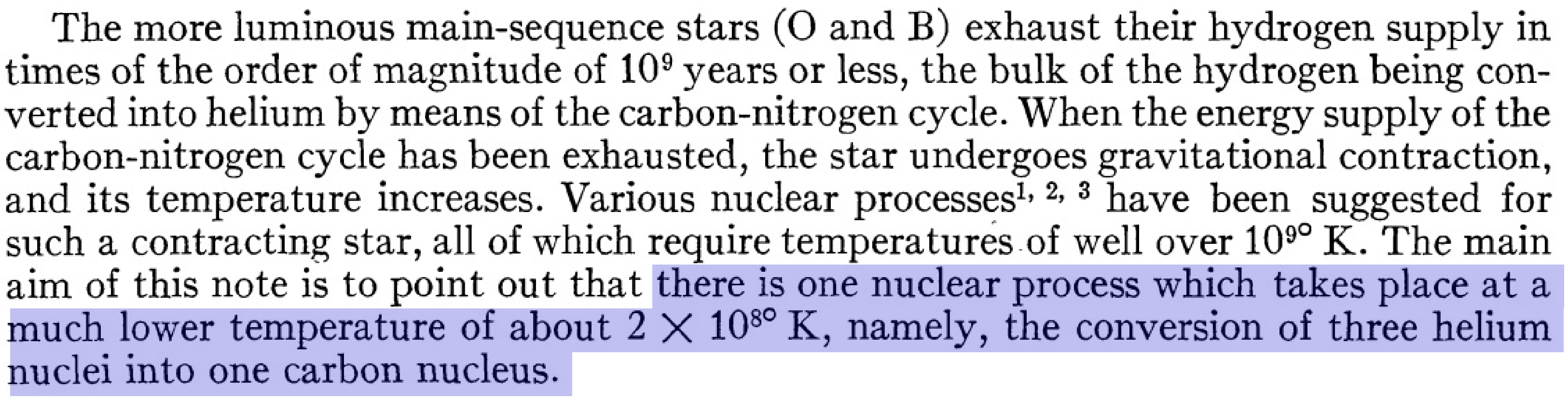
F. Hoyle:
all elements produced inside stars during their collapsing stage. by thermonuclear reactions



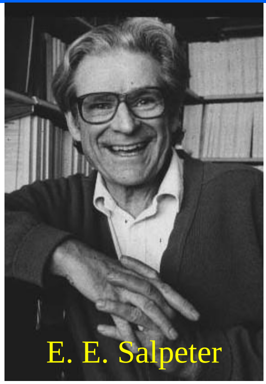
Old stars of galactic halo (population II) contain less heavy elements (metals) than younger stellar population (population I) of the galactic disc
Chamberlain and Aller 1951

The chemical composition of the Milky Way was substantially different in the past

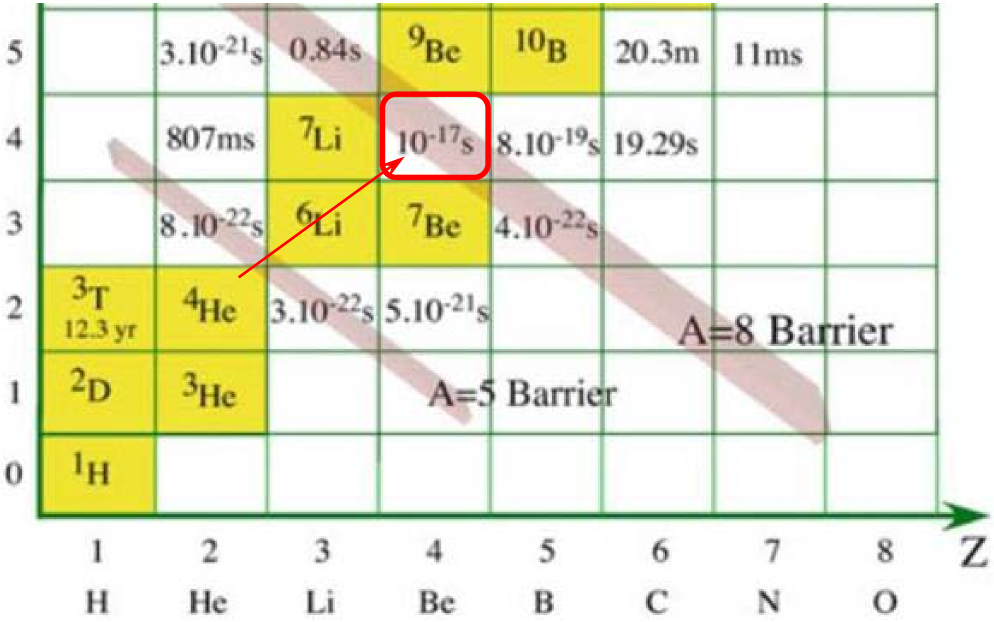
Nuclear reactions in stars without Hydrogen

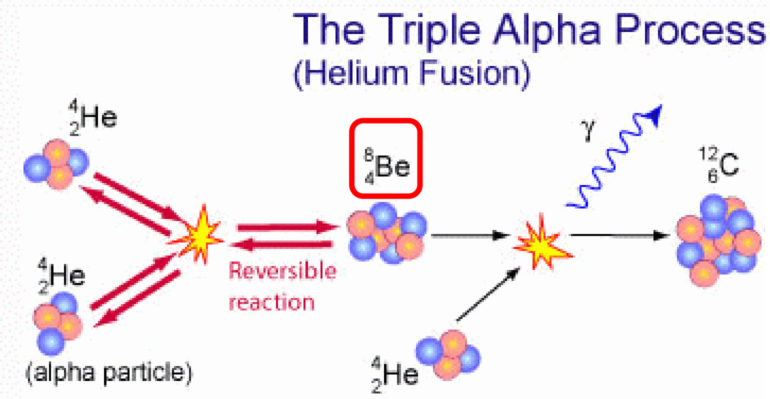


The more luminous main-sequence stars (O and B) exhaust their hydrogen supply in times of the order of magnitude of 10^9 years or less, the bulk of the hydrogen being converted into helium by means of the carbon-nitrogen cycle. When the energy supply of the carbon-nitrogen cycle has been exhausted, the star undergoes gravitational contraction, and its temperature increases. Various nuclear processes[1, 2, 3] have been suggested for such a contracting star, all of which require temperatures of well over 10^{9}° K. The main aim of this note is to point out that there is one nuclear process which takes place at a much lower temperature of about 2×10^{8}° K, namely, the conversion of three helium nuclei into one carbon nucleus.



E. E. Salpeter





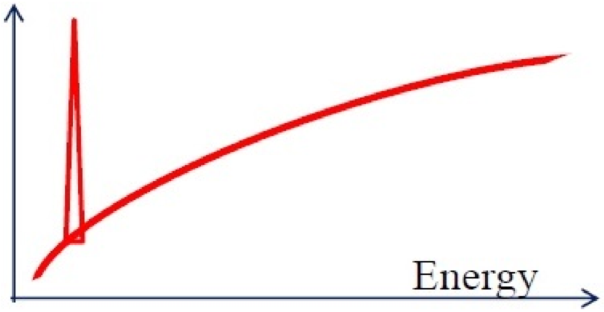
1st step
OK

2nd step
Not OK!

Formation and survival of ^{12}C in He-burning

Probability of nuclear reaction

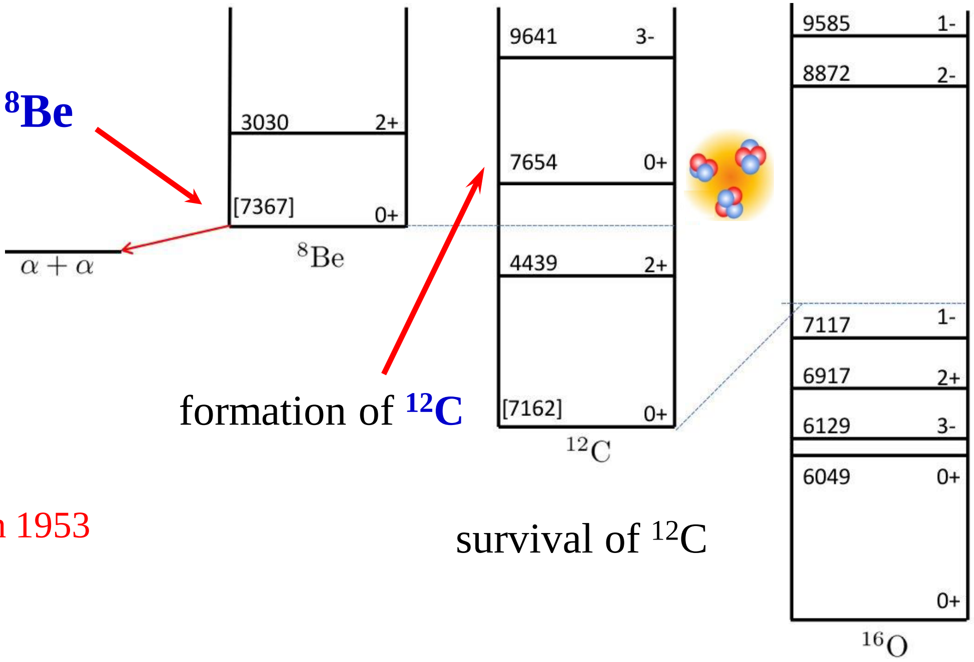
it becomes very high when the reaction is resonant



E. E. Salpeter

formation of ^{8}**Be**

Fred Hoyle suggests that the reaction $^{8}Be + \alpha \Longrightarrow {}^{12}C$ is resonant because of the existence of a nuclear energy level at 7.7 MeV (*unknown at that time*) in ^{12}C



formation of ^{12}**C**

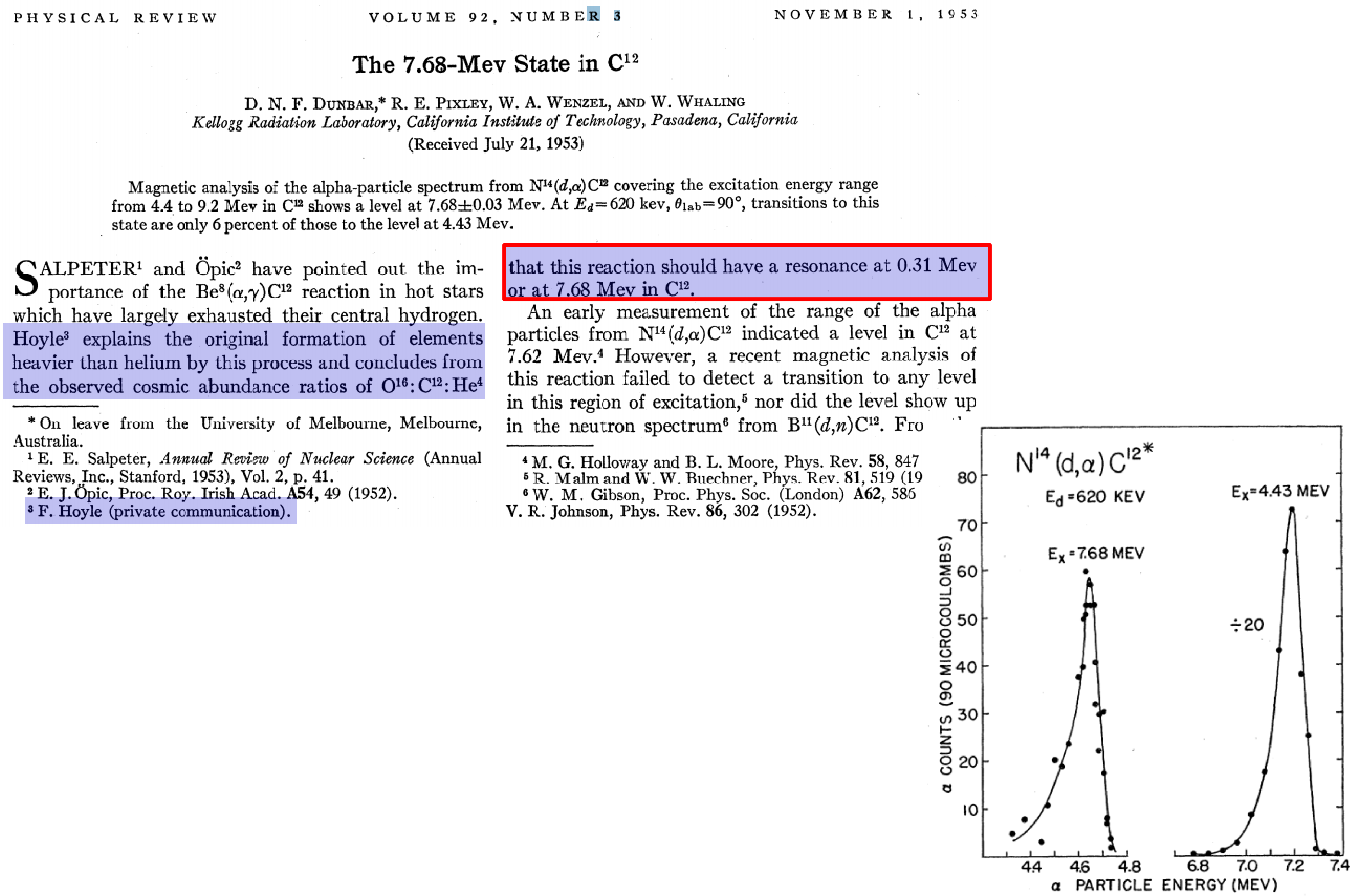
survival of ^{12}C

The level is found in W. Fowler´s Kellog laboratory in 1953

First quantitative prediction of a microscopic property of matter (structure of ^{12}C nucleus) from a macroscopic one (abundance of ^{12}C and ^{16}O)

1st and only ´prediction´ of the Anthropic Principle

Formation of ^{12}C



PHYSICAL REVIEW VOLUME 92, NUMBER 3 NOVEMBER 1, 1953

The 7.68-Mev State in C^{12}

D. N. F. Dunbar,* R. E. Pixley, W. A. Wenzel, and W. Whaling
Kellogg Radiation Laboratory, California Institute of Technology, Pasadena, California

(Received July 21, 1953)

Magnetic analysis of the alpha-particle spectrum from $N^{14}(d,\alpha)C^{12}$ covering the excitation energy range from 4.4 to 9.2 Mev in C^{12} shows a level at 7.68±0.03 Mev. At $E_d=620$ kev, $\theta_{lab}=90°$, transitions to this state are only 6 percent of those to the level at 4.43 Mev.

SALPETER[1] and Öpic[2] have pointed out the importance of the $Be^8(\alpha,\gamma)C^{12}$ reaction in hot stars which have largely exhausted their central hydrogen. Hoyle[3] explains the original formation of elements heavier than helium by this process and concludes from the observed cosmic abundance ratios of $O^{16}:C^{12}:He^4$ that this reaction should have a resonance at 0.31 Mev or at 7.68 Mev in C^{12}.

An early measurement of the range of the alpha particles from $N^{14}(d,\alpha)C^{12}$ indicated a level in C^{12} at 7.62 Mev.[4] However, a recent magnetic analysis of this reaction failed to detect a transition to any level in this region of excitation,[5] nor did the level show up in the neutron spectrum[6] from $B^{11}(d,n)C^{12}$. Fro

---

[*] On leave from the University of Melbourne, Melbourne, Australia.

[1] E. E. Salpeter, *Annual Review of Nuclear Science* (Annual Reviews, Inc., Stanford, 1953), Vol. 2, p. 41.

[2] E. J. Öpic, Proc. Roy. Irish Acad. A54, 49 (1952).

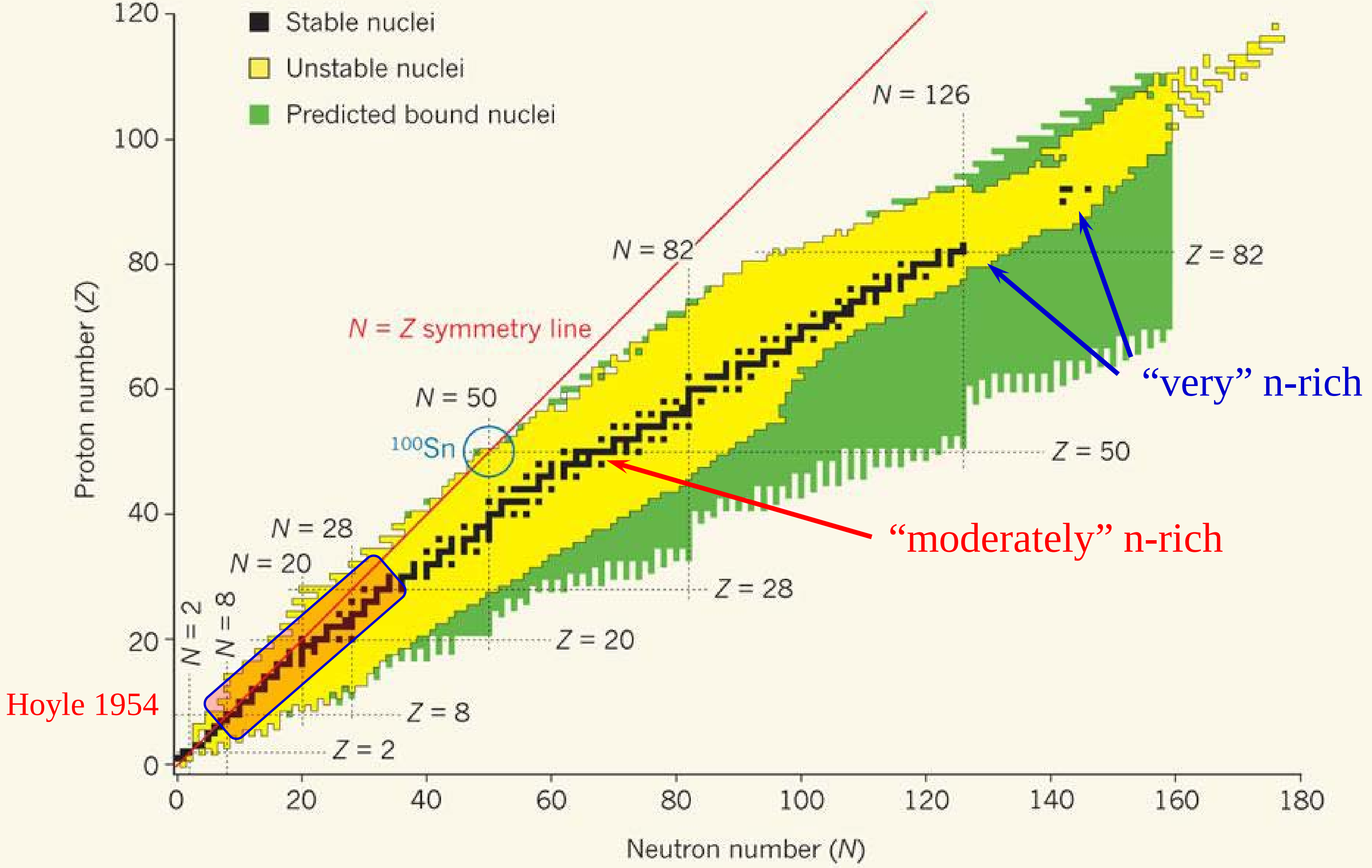
[3] F. Hoyle (private communication).

[4] M. G. Holloway and B. L. Moore, Phys. Rev. **58**, 847

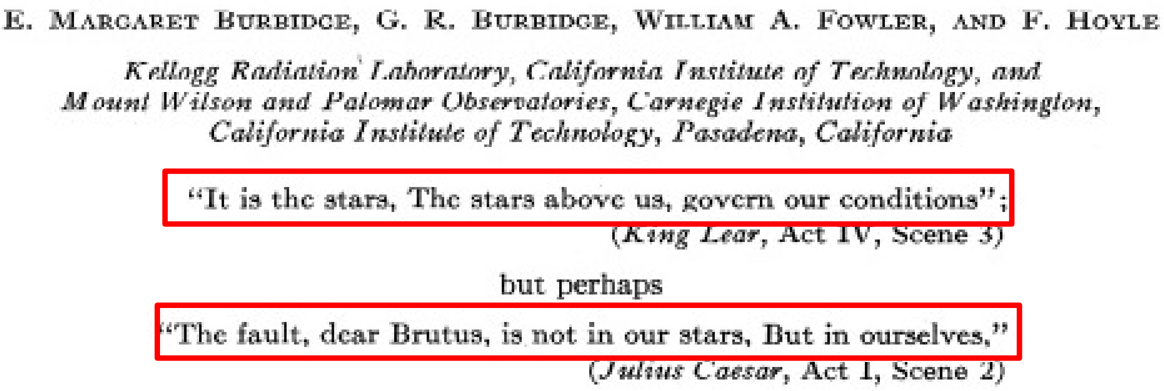
[5] R. Malm and W. W. Buechner, Phys. Rev. **81**, 519 (19

[6] W. M. Gibson, Proc. Phys. Soc. (London) A62, 586 V. R. Johnson, Phys. Rev. **86**, 302 (1952).

Elements and Isotopes



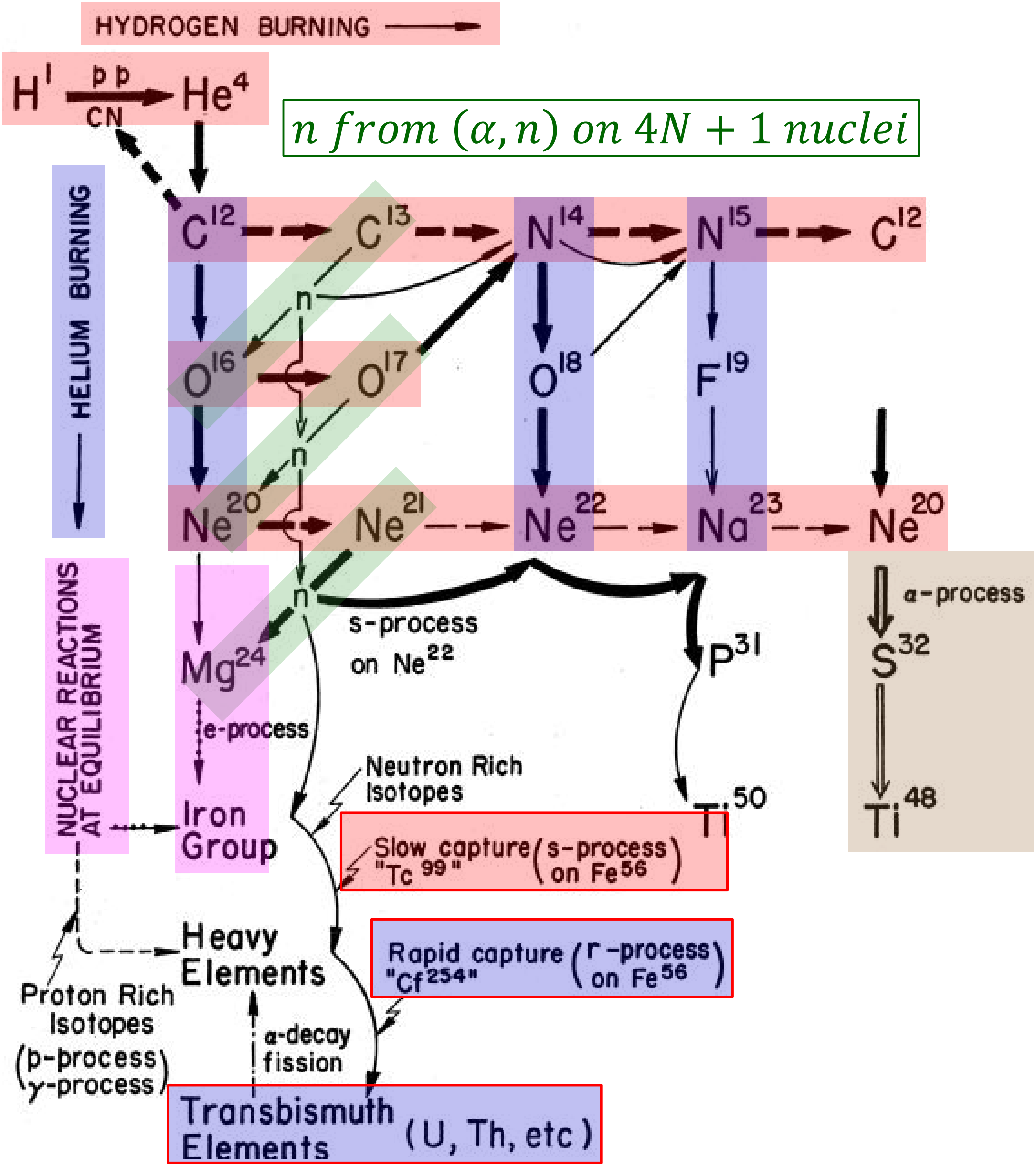
Synthesis of the elements in stars 1957



E. MARGARET BURBIDGE, G. R. BURBIDGE, WILLIAM A. FOWLER, AND F. HOYLE

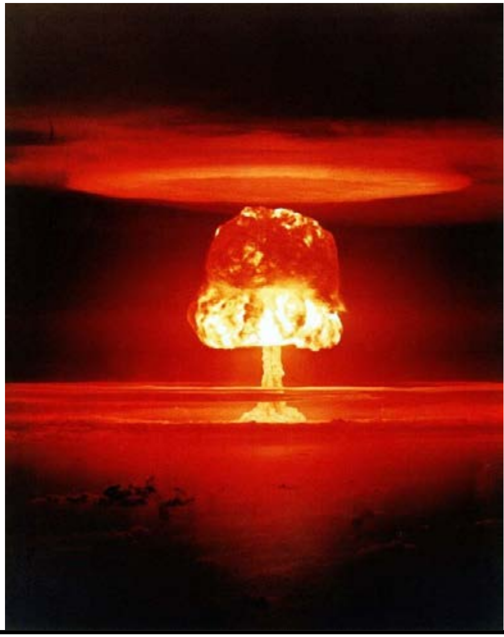
Kellogg Radiation Laboratory, California Institute of Technology, and
Mount Wilson and Palomar Observatories, Carnegie Institution of Washington,
California Institute of Technology, Pasadena, California

> "It is the stars, The stars above us, govern our conditions";
> (*King Lear*, Act IV, Scene 3)
>
> but perhaps
>
> "The fault, dear Brutus, is not in our stars, But in ourselves,"
> (*Julius Caesar*, Act I, Scene 2)



$n\ from\ (\alpha,n)\ on\ 4N+1\ nuclei$

Data on neutron capture cross sections and yields from the first H-bomb test in Bikini island (1952)



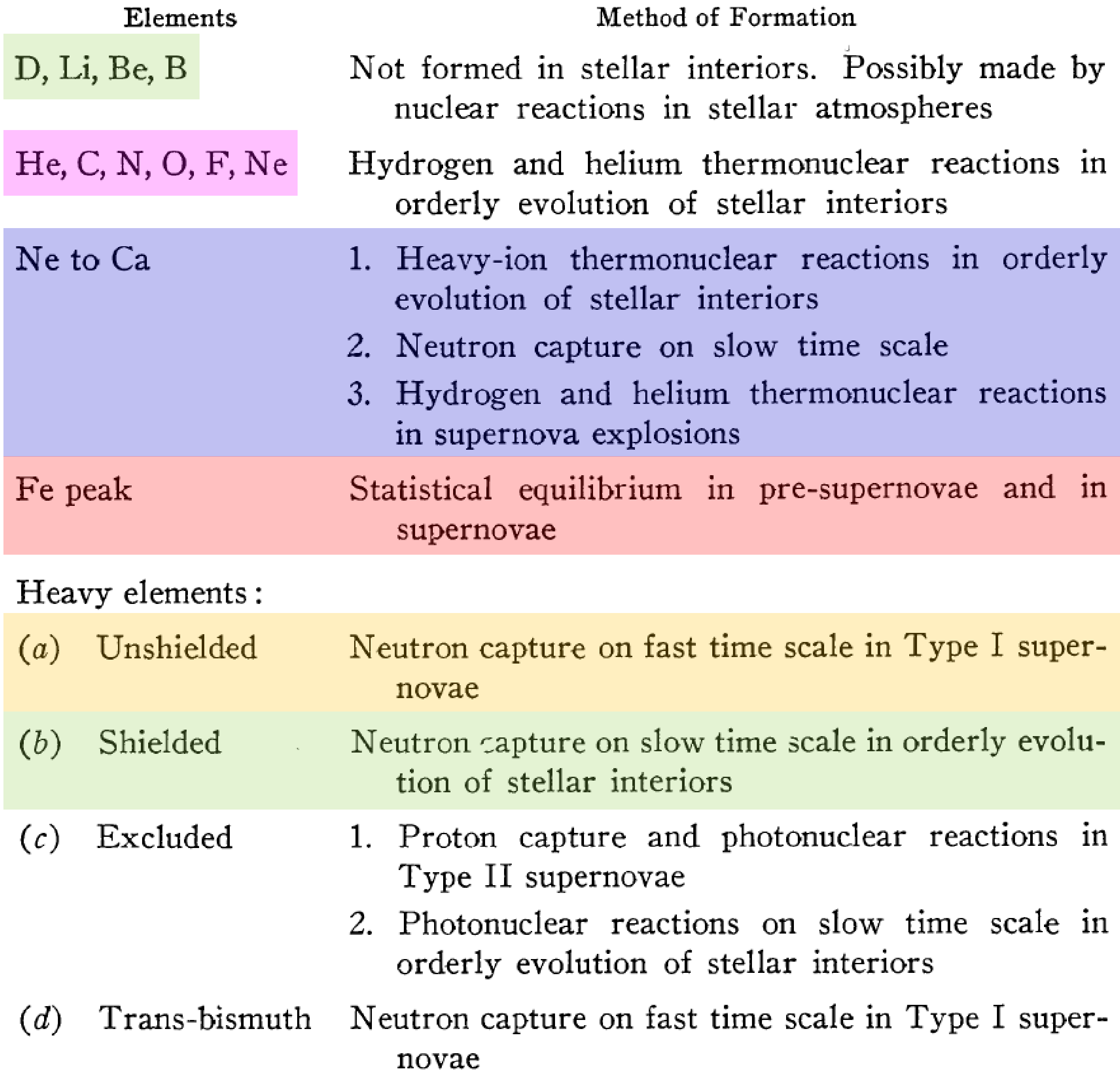
Method of formation

1957:

Alastair G.W. Cameron

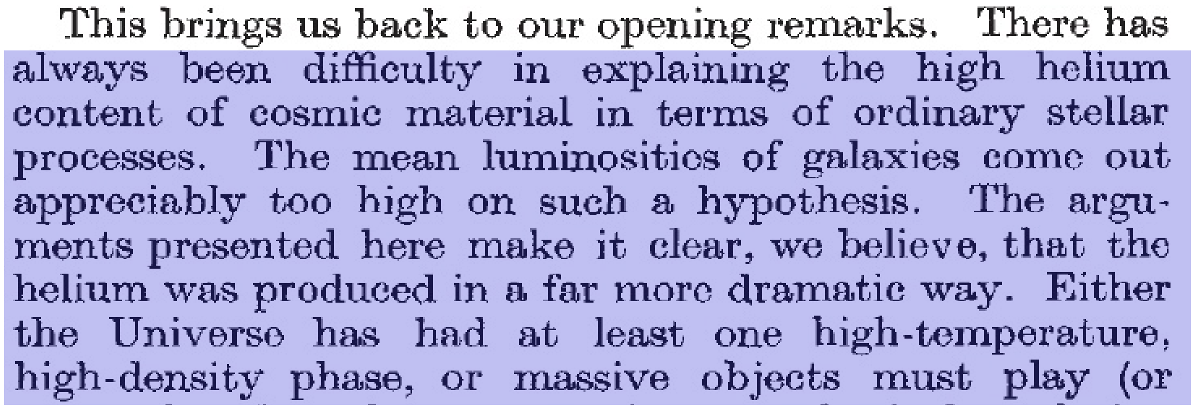
Nuclear reactions
in stars and nucleogenesis

(Chalk River report)

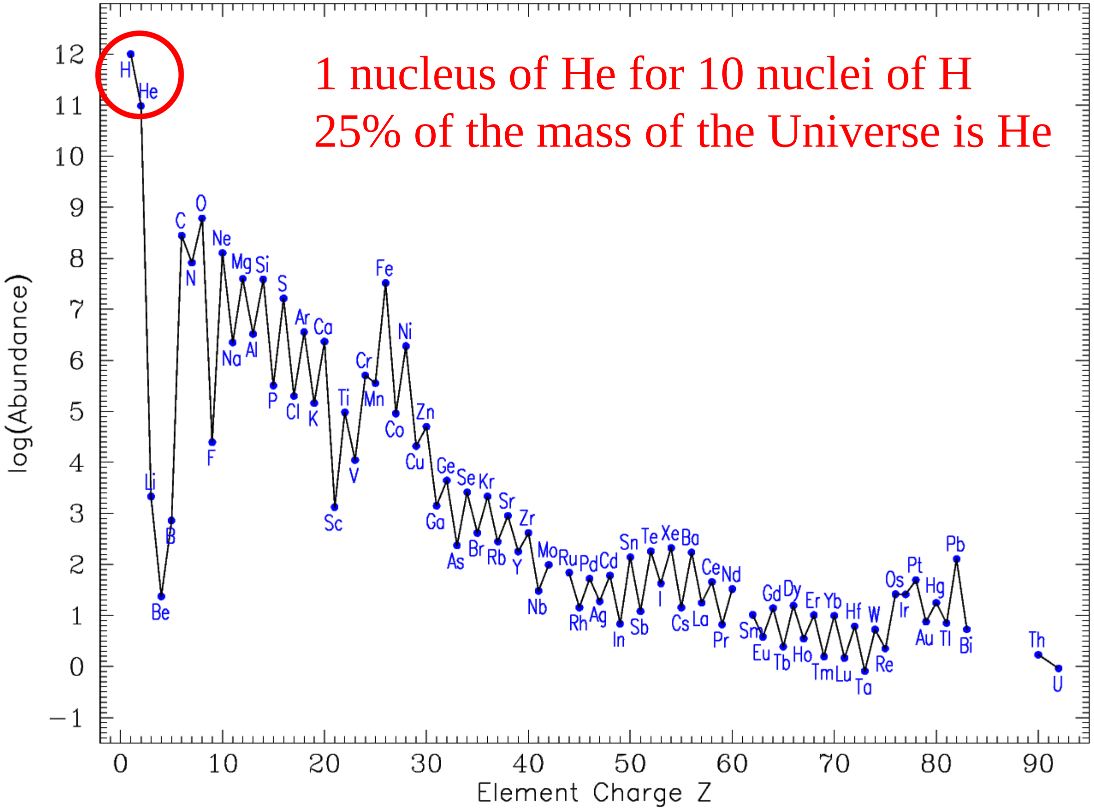


Elements	Method of Formation
D, Li, Be, B	Not formed in stellar interiors. Possibly made by nuclear reactions in stellar atmospheres
He, C, N, O, F, Ne	Hydrogen and helium thermonuclear reactions in orderly evolution of stellar interiors
Ne to Ca	1. Heavy-ion thermonuclear reactions in orderly evolution of stellar interiors 2. Neutron capture on slow time scale 3. Hydrogen and helium thermonuclear reactions in supernova explosions
Fe peak	Statistical equilibrium in pre-supernovae and in supernovae
Heavy elements:	
(*a*) Unshielded	Neutron capture on fast time scale in Type I supernovae
(*b*) Shielded	Neutron capture on slow time scale in orderly evolution of stellar interiors
(*c*) Excluded	1. Proton capture and photonuclear reactions in Type II supernovae 2. Photonuclear reactions on slow time scale in orderly evolution of stellar interiors
(*d*) Trans-bismuth	Neutron capture on fast time scale in Type I supernovae

The mystery of the cosmic Helium abundance 1964

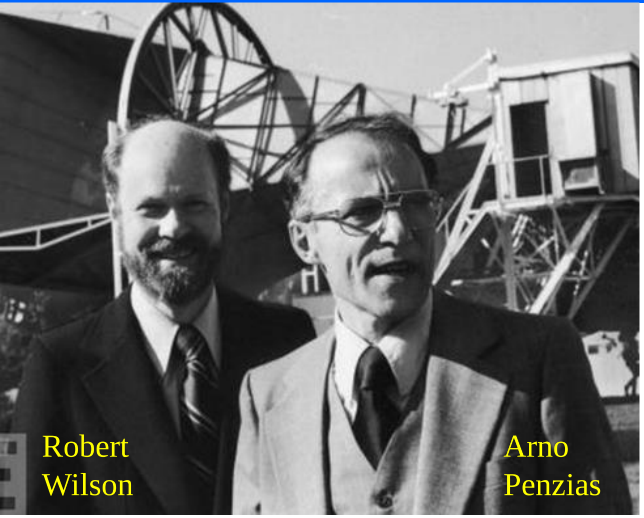


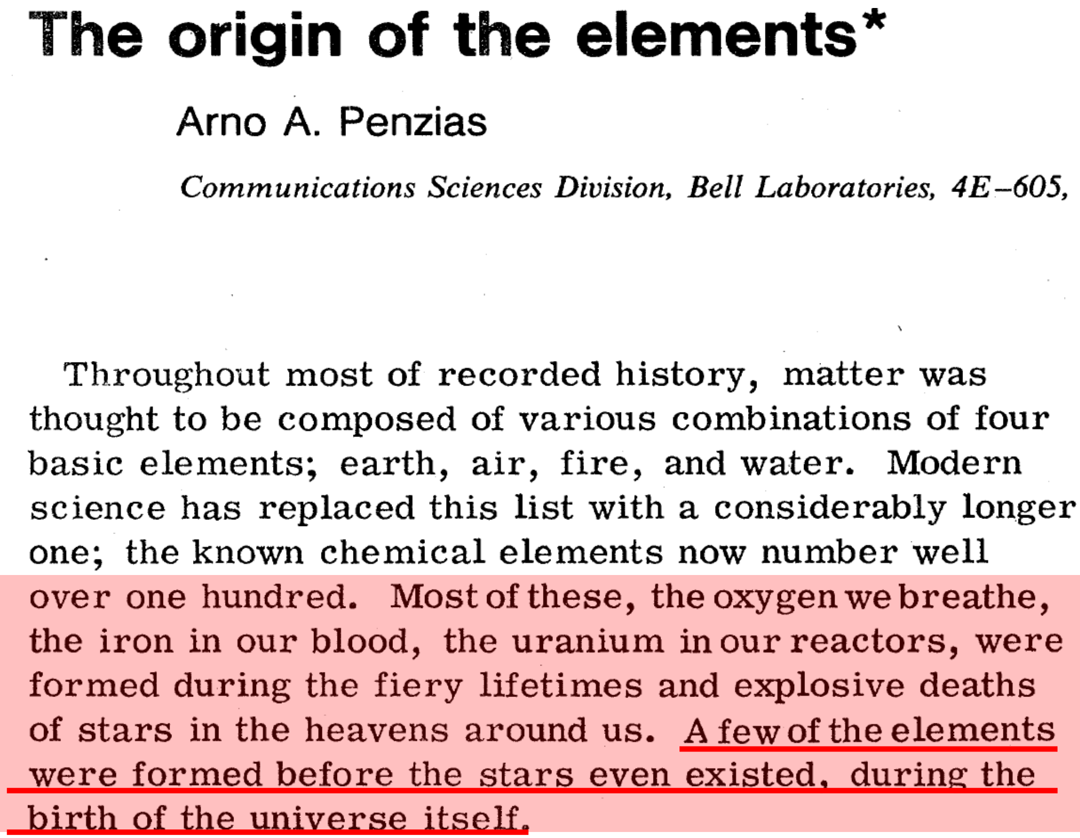
This brings us back to our opening remarks. There has always been difficulty in explaining the high helium content of cosmic material in terms of ordinary stellar processes. The mean luminosities of galaxies come out appreciably too high on such a hypothesis. The arguments presented here make it clear, we believe, that the helium was produced in a far more dramatic way. Either the Universe has had at least one high-temperature, high-density phase, or massive objects must play (or



1 nucleus of He for 10 nuclei of H
25% of the mass of the Universe is He

Discovery of the Cosmic microwave background 1965



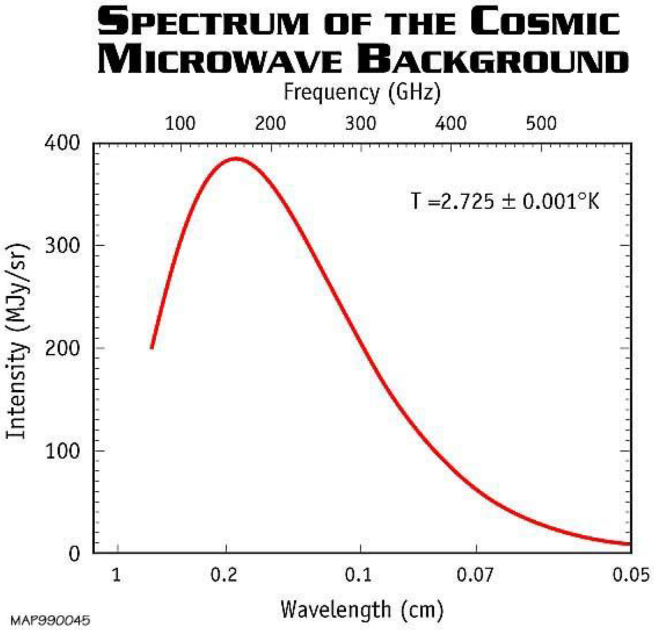


The origin of the elements*

Arno A. Penzias

Communications Sciences Division, Bell Laboratories, 4E–605,

Throughout most of recorded history, matter was thought to be composed of various combinations of four basic elements; earth, air, fire, and water. Modern science has replaced this list with a considerably longer one; the known chemical elements now number well over one hundred. Most of these, the oxygen we breathe, the iron in our blood, the uranium in our reactors, were formed during the fiery lifetimes and explosive deaths of stars in the heavens around us. A few of the elements were formed before the stars even existed, during the birth of the universe itself.



On the synthesis of elements at very high temperatures 1966

R. V. Wagoner, W. A. Fowler, F. Hoyle

A detailed calculation of element production in the early stages of a homogeneous and isotropic expanding universe as well as within imploding-exploding supermassive stars has been made. If the recently measured microwave background radiation is due to primeval photons, then significant quantities of only D, He^3, He^4, and Li^7 can be produced in the universal fireball. Reasonable agreement with solar-system abundances for these nuclei is obtained if the present temperature is 3° K and if the present density is $\sim 2 \times 10^{-31}$ gm cm^3, corresponding to a deceleration parameter $q_0 \approx 5 \times 10^{-3}$. However,

Hayashi (1950): n – p equilibrium

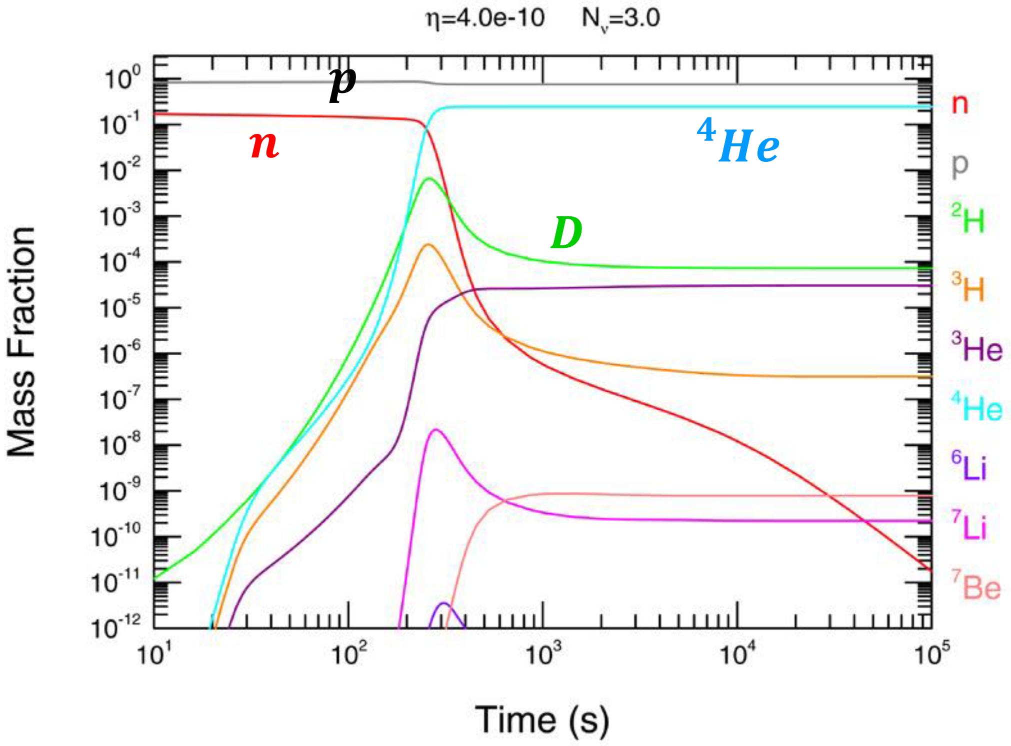
$X(n)/X(p) \sim 0.13$ at freeze-out

$X(^4He) \sim 2 \cdot X(n) \sim 0.25$

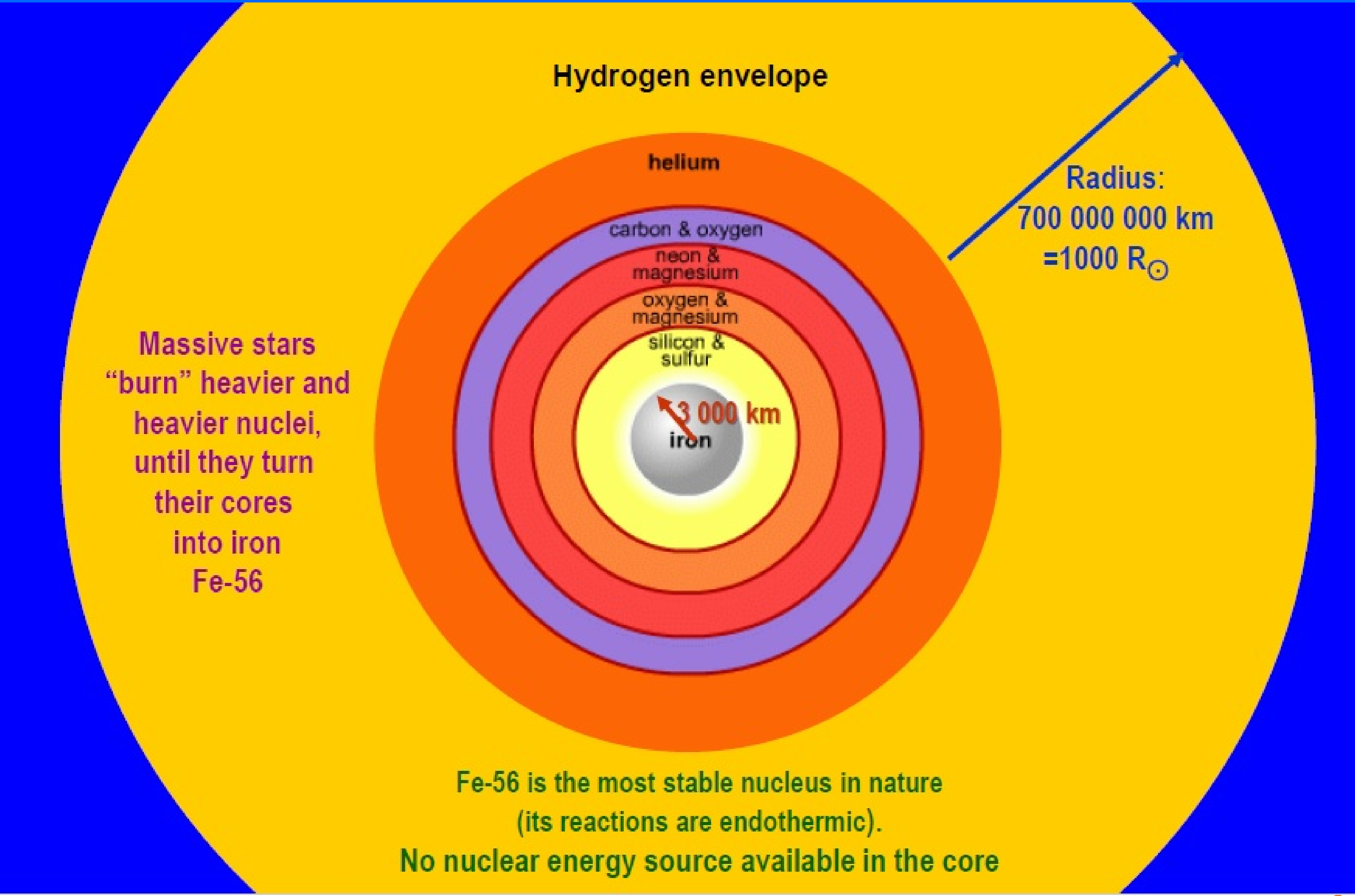
only way to produce so much ^4He ~ (25% by mass)

only way to produce Deuterium ~ (destroyed in stellar interiors)

in excellent agreement with observation



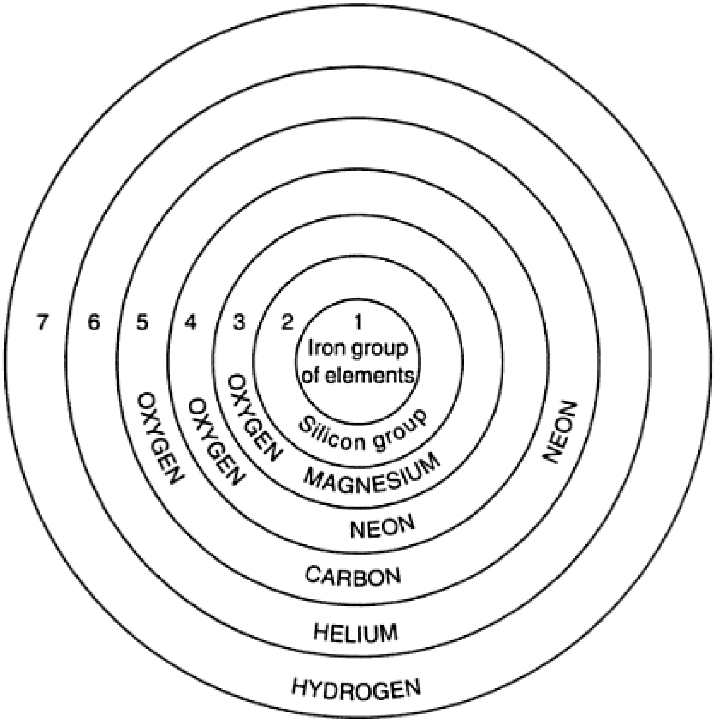
Back to the stars



Massive stars
“burn” heavier and
heavier nuclei,
until they turn
their cores
into iron
Fe-56

Fe-56 is the most stable nucleus in nature
(its reactions are endothermic).
No nuclear energy source available in the core

Advanced nucleosynthesis phases in massive stars



Because of the increased sensitivity of nuclear reactions to temperature heavier nuclei are produced closer to the center of the stars

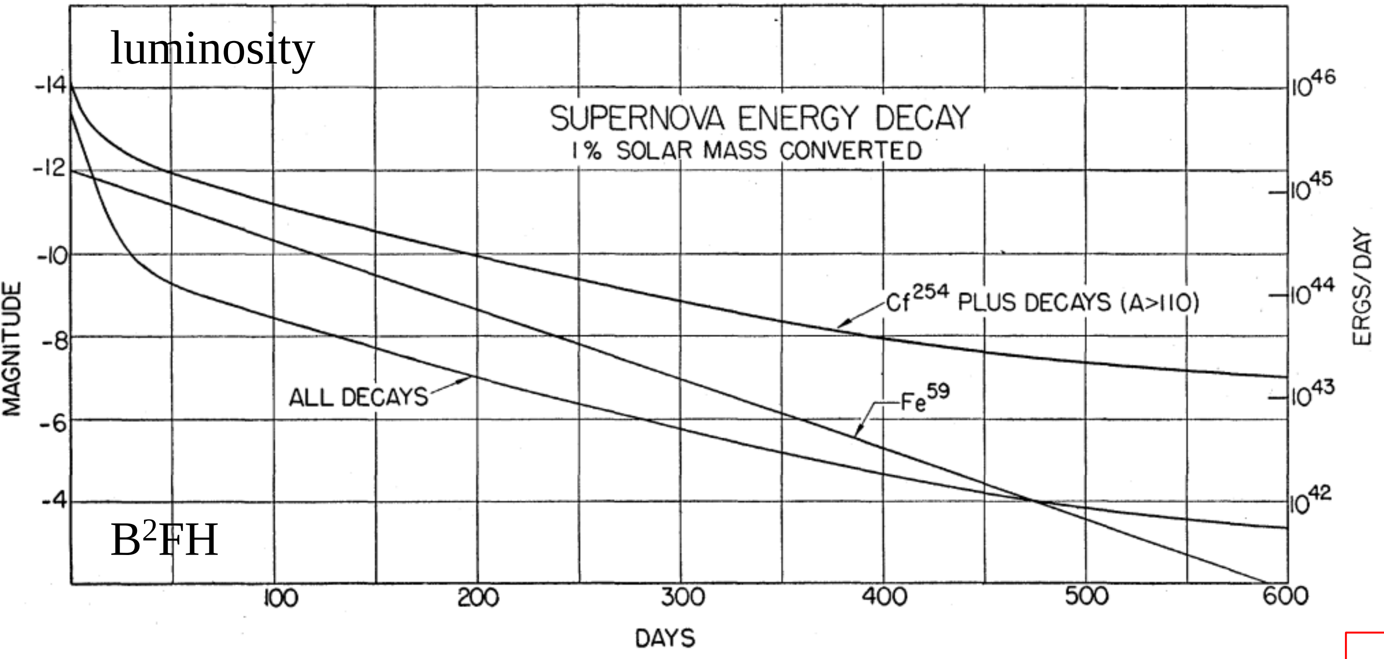
The “onion skin” model (Hoyle 1955)

The outer layers keep (partially) the products of previous phases of stellar nucleosynthesis

This material synthesized by nuclear reactions in stellar interiors must come out through the final event of a **supernova explosion**

The explosion produces also new elements (explosive nucleosynthesis) including short-lived **radioactive** ones

What powers the exponentially decreasing lightcurve of supernovae?



Radioactivity

of lifetime ~ 2 months

^{7}Be: Borst 1950

^{254}Cf: Baade, B2FH 1956

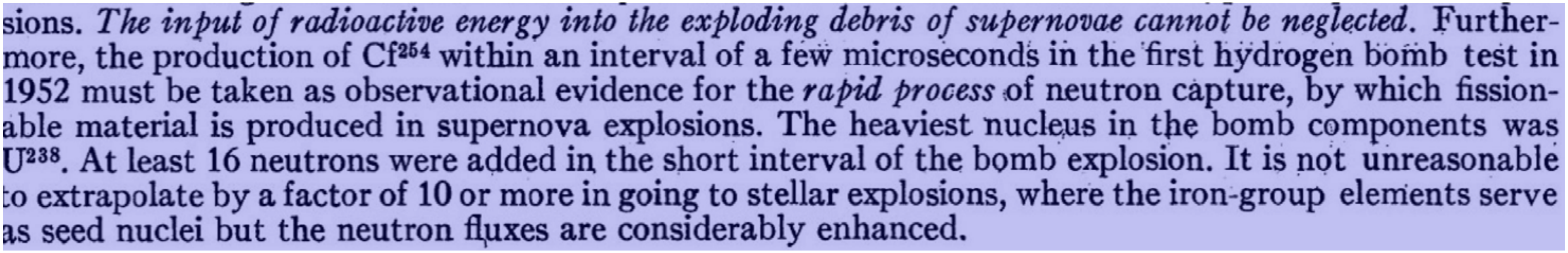
$^{56}\text{Ni} \rightarrow {}^{56}\text{Co} \rightarrow {}^{56}\text{Fe}$

7 days *2 months*

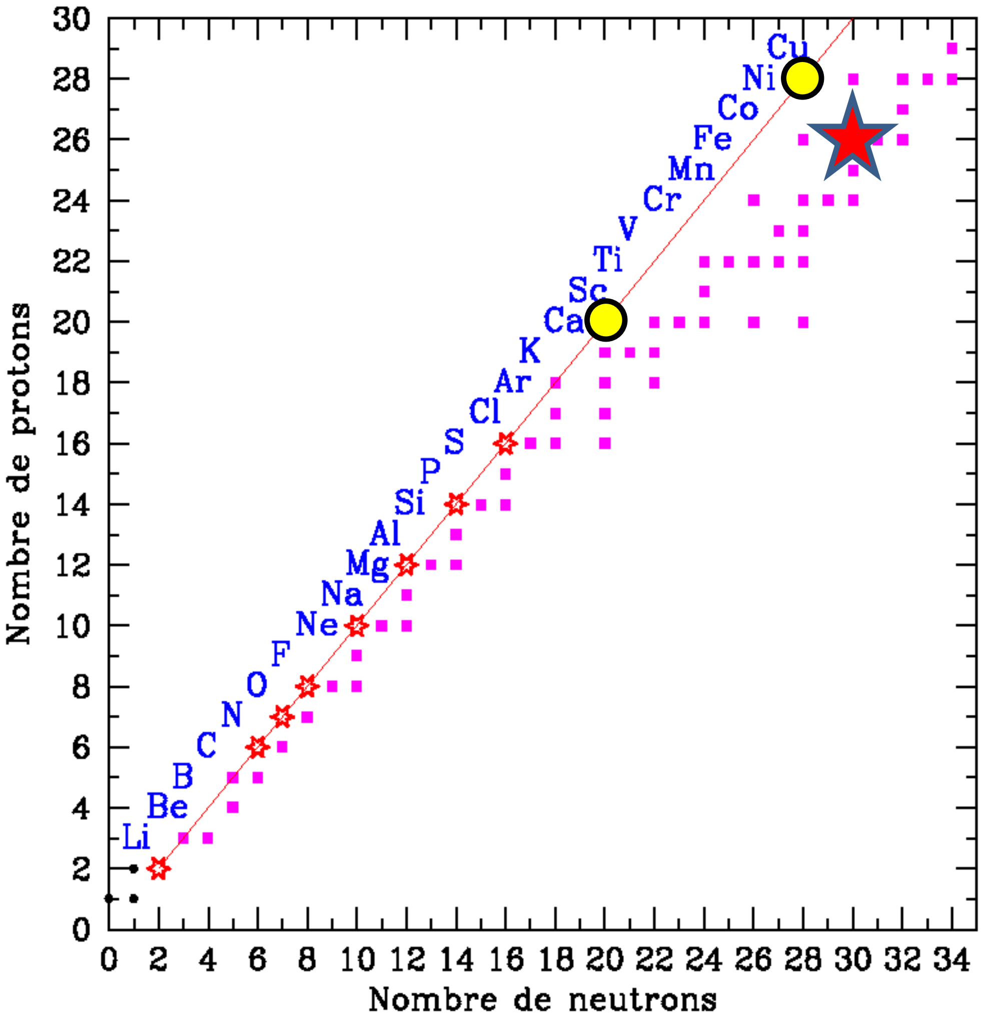
^{56}Fe, the most stable nucleus in nature is produced as unstable ^{56}Ni

Hoyle´s greatest regret

Hoyle and Fowler 1960: Explosive nucleosynthesis


sions. *The input of radioactive energy into the exploding debris of supernovae cannot be neglected.* Furthermore, the production of Cf^{254} within an interval of a few microseconds in the first hydrogen bomb test in 1952 must be taken as observational evidence for the *rapid process* of neutron capture, by which fissionable material is produced in supernova explosions. The heaviest nucleus in the bomb components was U^{238}. At least 16 neutrons were added in the short interval of the bomb explosion. It is not unreasonable to extrapolate by a factor of 10 or more in going to stellar explosions, where the iron-group elements serve as seed nuclei but the neutron fluxes are considerably enhanced.

The way towards the Fe-peak: hydrostatic versus explosive



^{40}Ca is the last stable nucleus with N=Z on the way of Si-melting towards the Fe-peak

In the stellar core **weak interactions**

$$p \rightarrow n + e^{+} + \nu$$

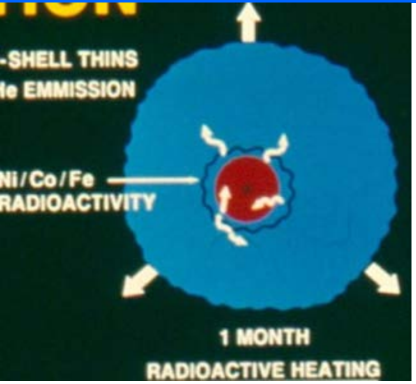
turn some protons into neutrons and ^{56}Fe dominates the composition in nuclear statistical equilibrium

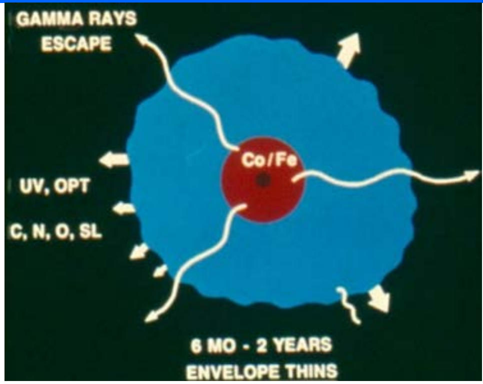
In explosive nucleosynthesis, weak (= slow) interactions have no time to operate and the nuclear flow goes through N = Z up to ^{56}Ni (N=Z=28)

Hoyle´s greatest regret: missing the origin of ^{56}Fe, the most stable nucleus in nature

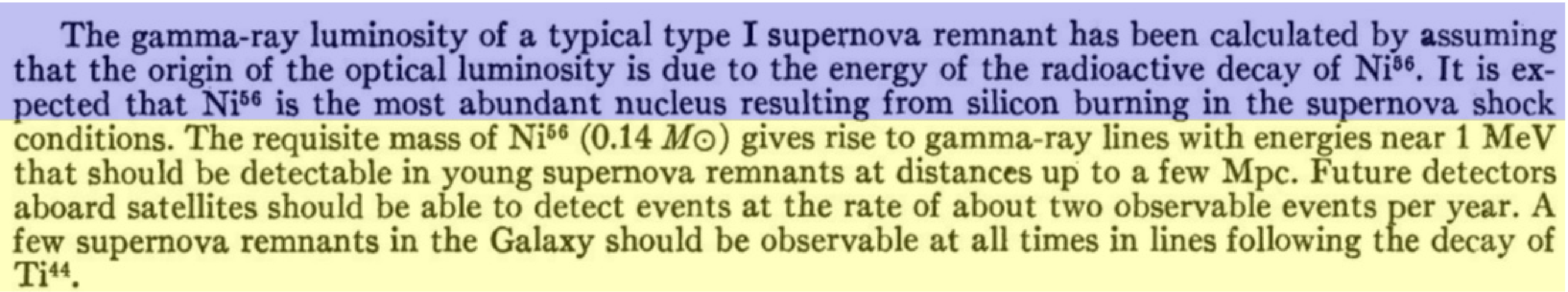
It is produced as unstable ^{56}Ni in supernova explosions

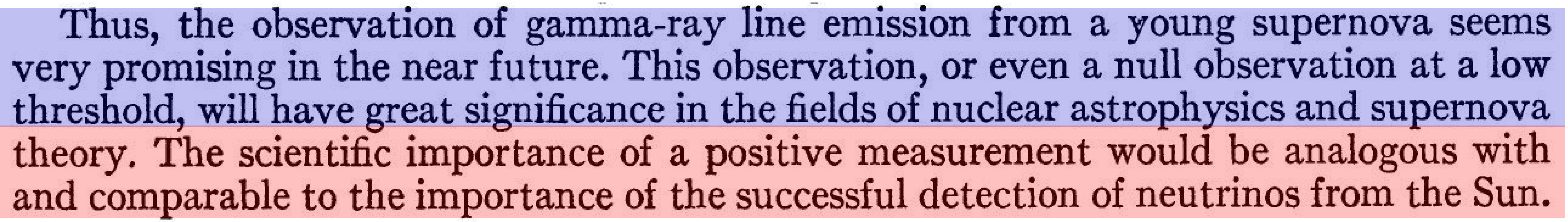
Gamma-ray lines from young supernova remnants





D. D. Clayton, S. A. Colgate, G. J. Fishman (1968)


The gamma-ray luminosity of a typical type I supernova remnant has been calculated by assuming that the origin of the optical luminosity is due to the energy of the radioactive decay of Ni^{56}. It is expected that Ni^{56} is the most abundant nucleus resulting from silicon burning in the supernova shock conditions. The requisite mass of Ni^{56} (0.14 $M\odot$) gives rise to gamma-ray lines with energies near 1 MeV that should be detectable in young supernova remnants at distances up to a few Mpc. Future detectors aboard satellites should be able to detect events at the rate of about two observable events per year. A few supernova remnants in the Galaxy should be observable at all times in lines following the decay of Ti^{44}.


Thus, the observation of gamma-ray line emission from a young supernova seems very promising in the near future. This observation, or even a null observation at a low threshold, will have great significance in the fields of nuclear astrophysics and supernova theory. The scientific importance of a positive measurement would be analogous with and comparable to the importance of the successful detection of neutrinos from the Sun.

The hydrodynamic behavior of supernovae explosion


Stirling Colgate

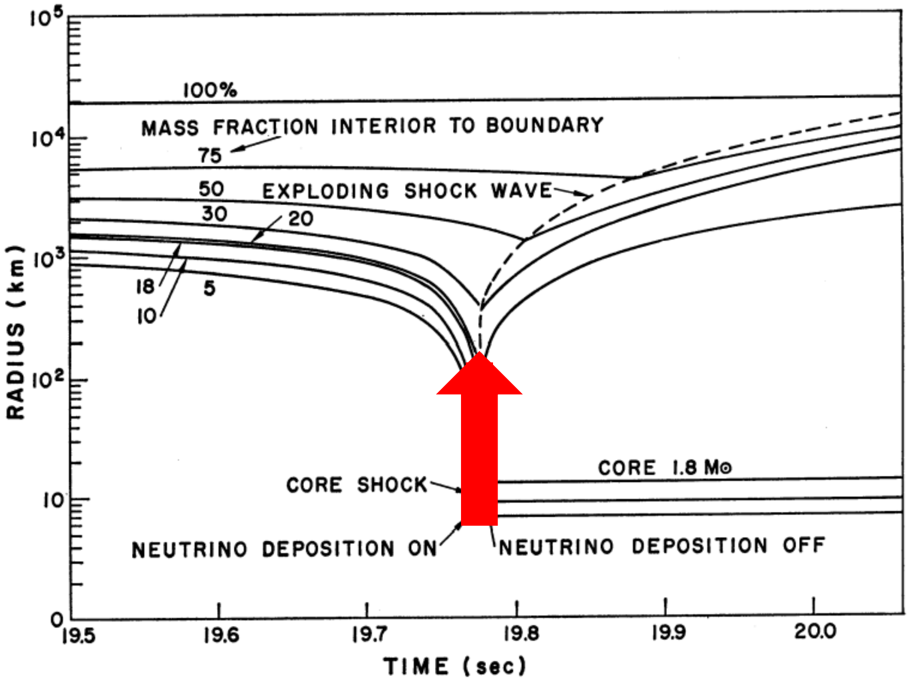
1959: suggests satellites for search of γ – rays from supernovae.

In fact to cover monitoring of soviet test explosions

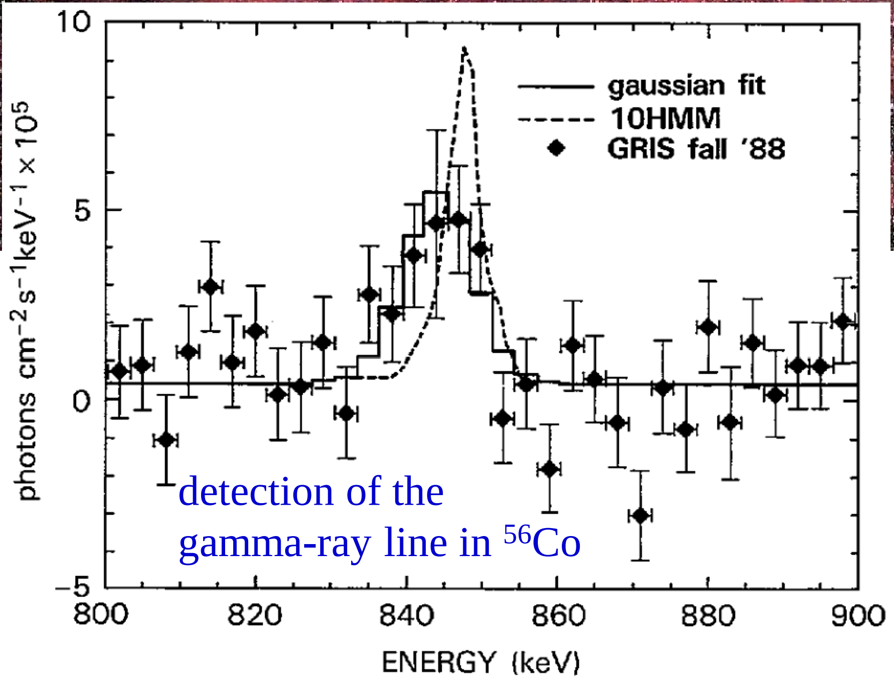
> We regard the release of gravitational energy attending a dynamic change in configuration to be the primary energy source in supernovae explosions. Although we were initially inspired by and agree in detail with the mechanism for initiating gravitational instability proposed by Burbidge, Burbidge, Fowler, and Hoyle, we find that the dynamical implosion is so violent that an energy many times greater than the available thermonuclear energy is released from the star's core and transferred to the star's mantle in a supernova explosion. The energy released corresponds to the change in gravitational potential of the unstable imploding core; the transfer of energy takes place by the emission and deposition of neutrinos.

99% of the enormous amount of gravitational energy from the dynamical collapse of the Fe core is released in the form of **neutrinos**

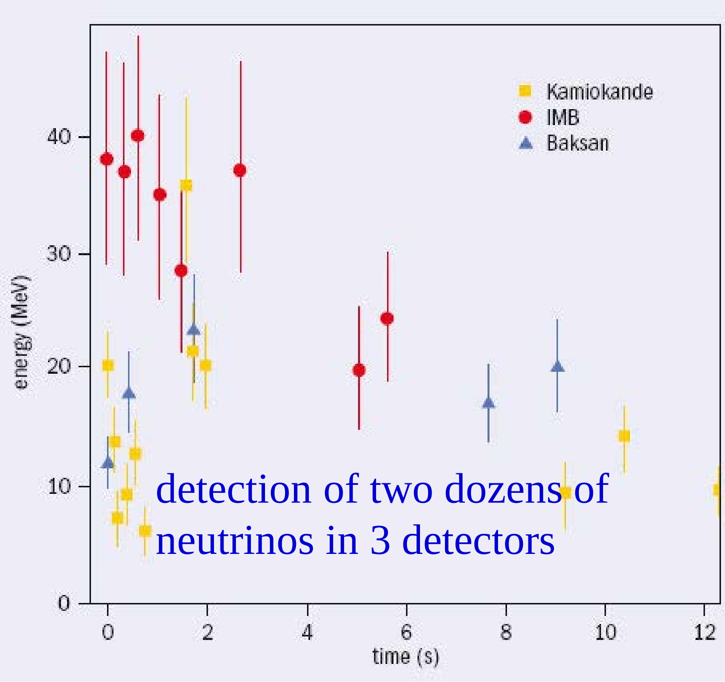
Some of them interact with the stellar mantle and expel it, making a supernova



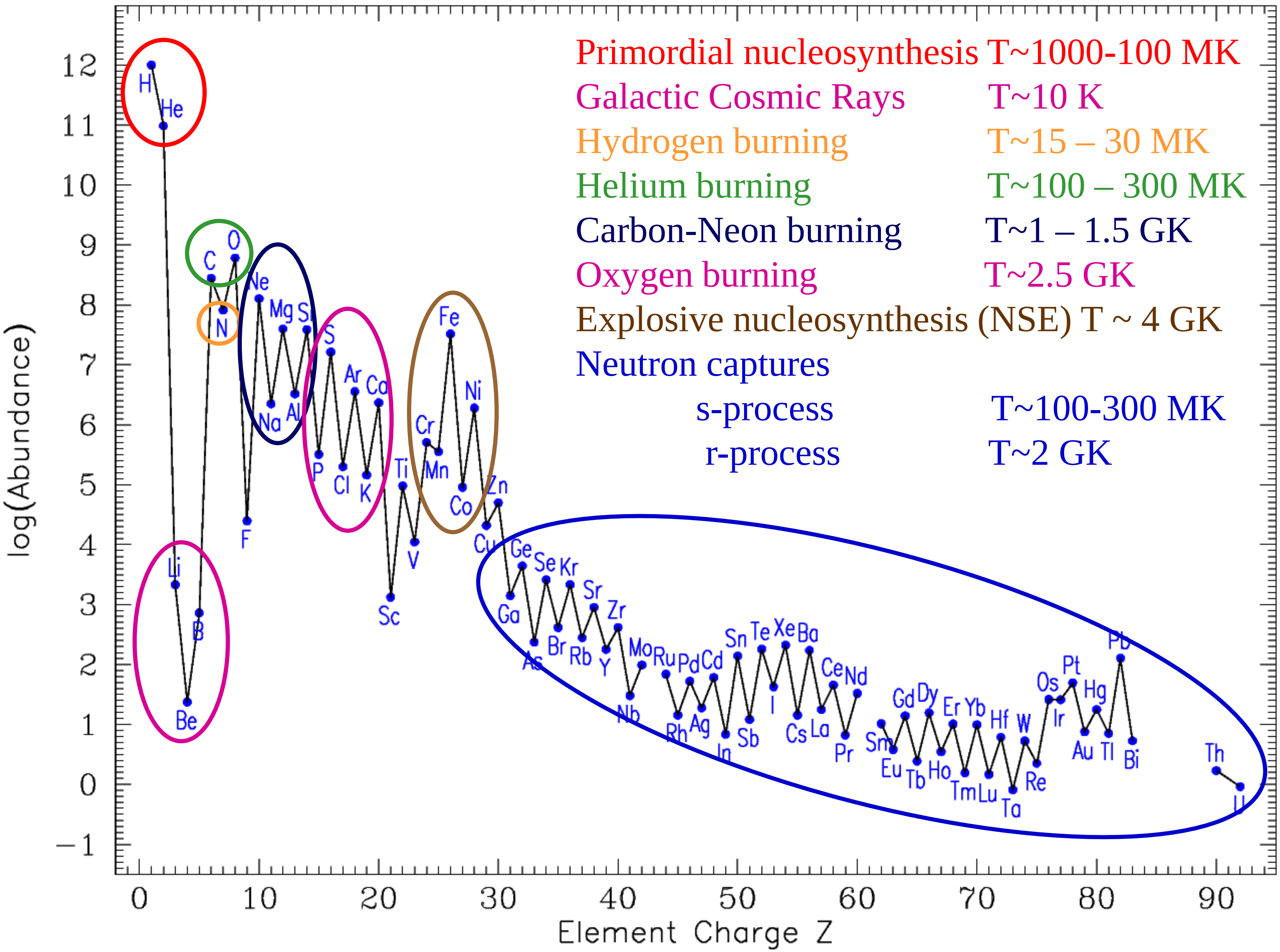
SN 1987A in the
Large Magellanic Cloud
150 000 light years from Earth



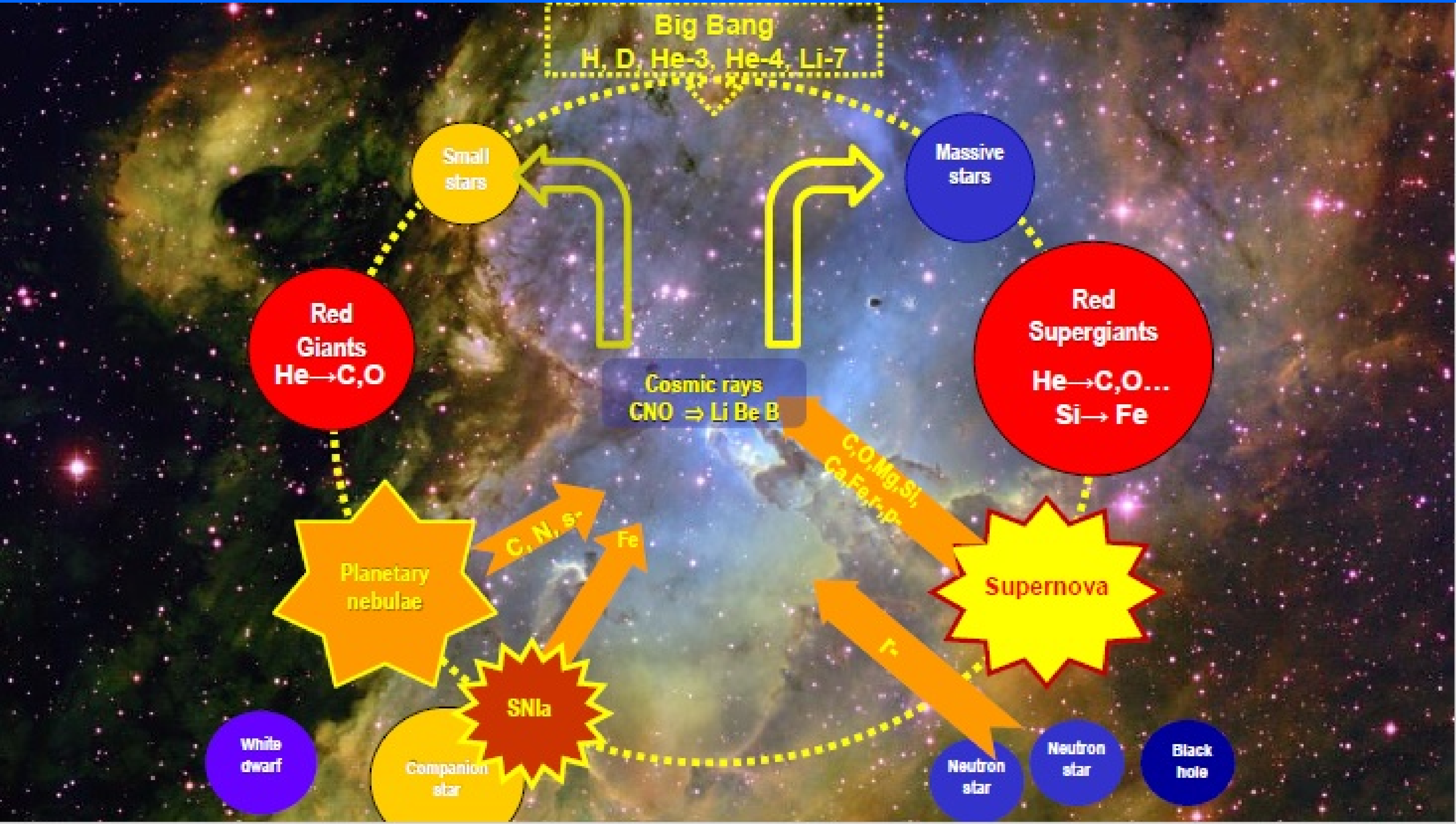
detection of the
gamma-ray line in ^{56}Co



detection of two dozens of
neutrinos in 3 detectors



Primordial nucleosynthesis	T~1000-100 MK
Galactic Cosmic Rays	T~10 K
Hydrogen burning	T~15 – 30 MK
Helium burning	T~100 – 300 MK
Carbon-Neon burning	T~1 – 1.5 GK
Oxygen burning	T~2.5 GK
Explosive nucleosynthesis (NSE) T ~ 4 GK	
Neutron captures	
s-process	T~100-300 MK
r-process	T~2 GK



Big Bang
H, D, He-3, He-4, Li-7

Small stars

Massive stars

Red Giants
He→C,O

Red Supergiants
He→C,O…
Si→ Fe

Cosmic rays
CNO ⇒ Li Be B

C,O,Mg,Si,
Ca,Fe,r,p-

C, N, s-

Fe

Planetary nebulae

Supernova

r-

SNIa

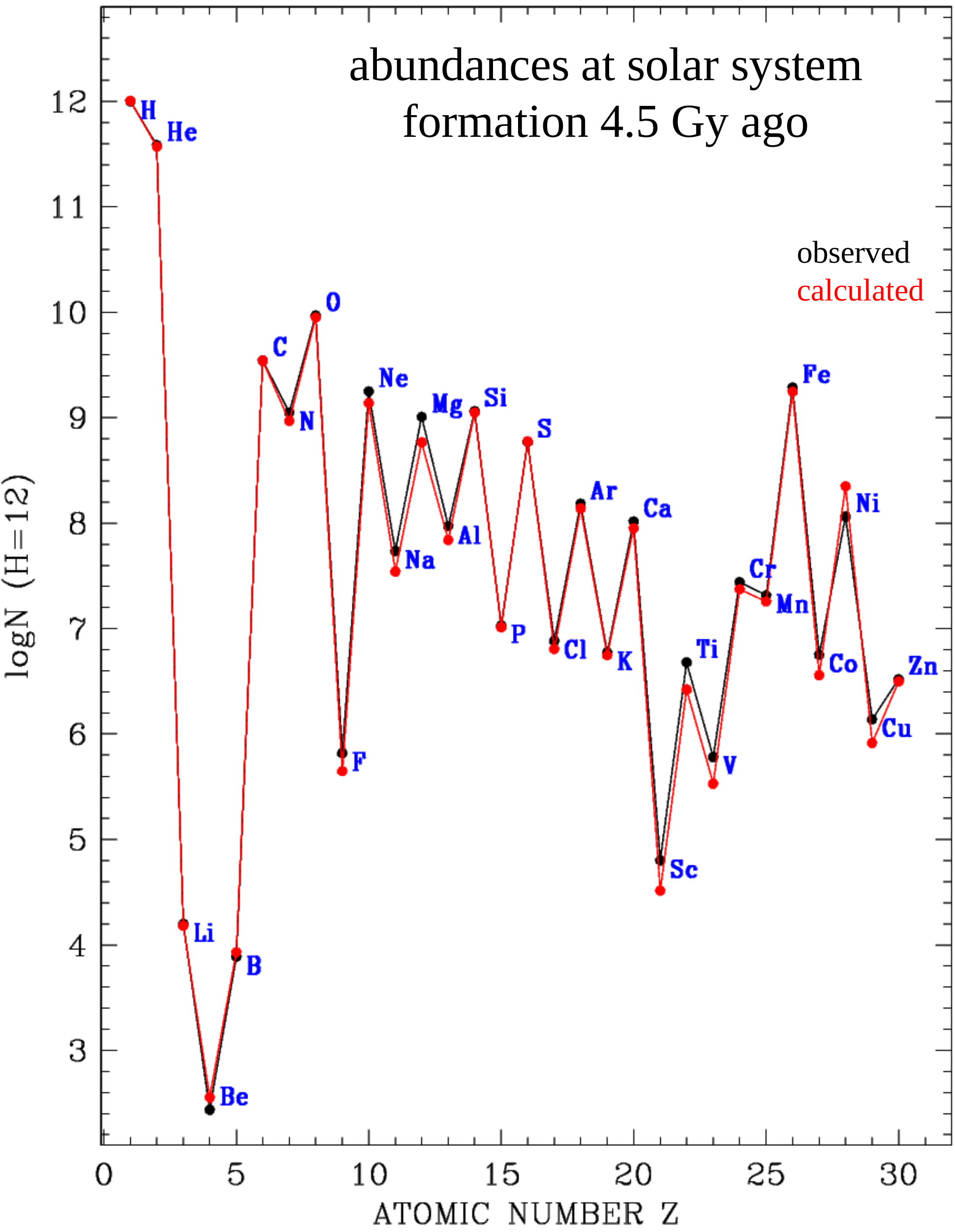
White dwarf

Companion star

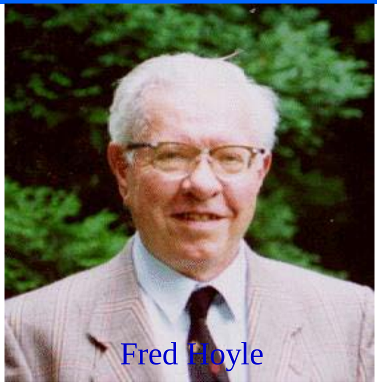
Neutron star

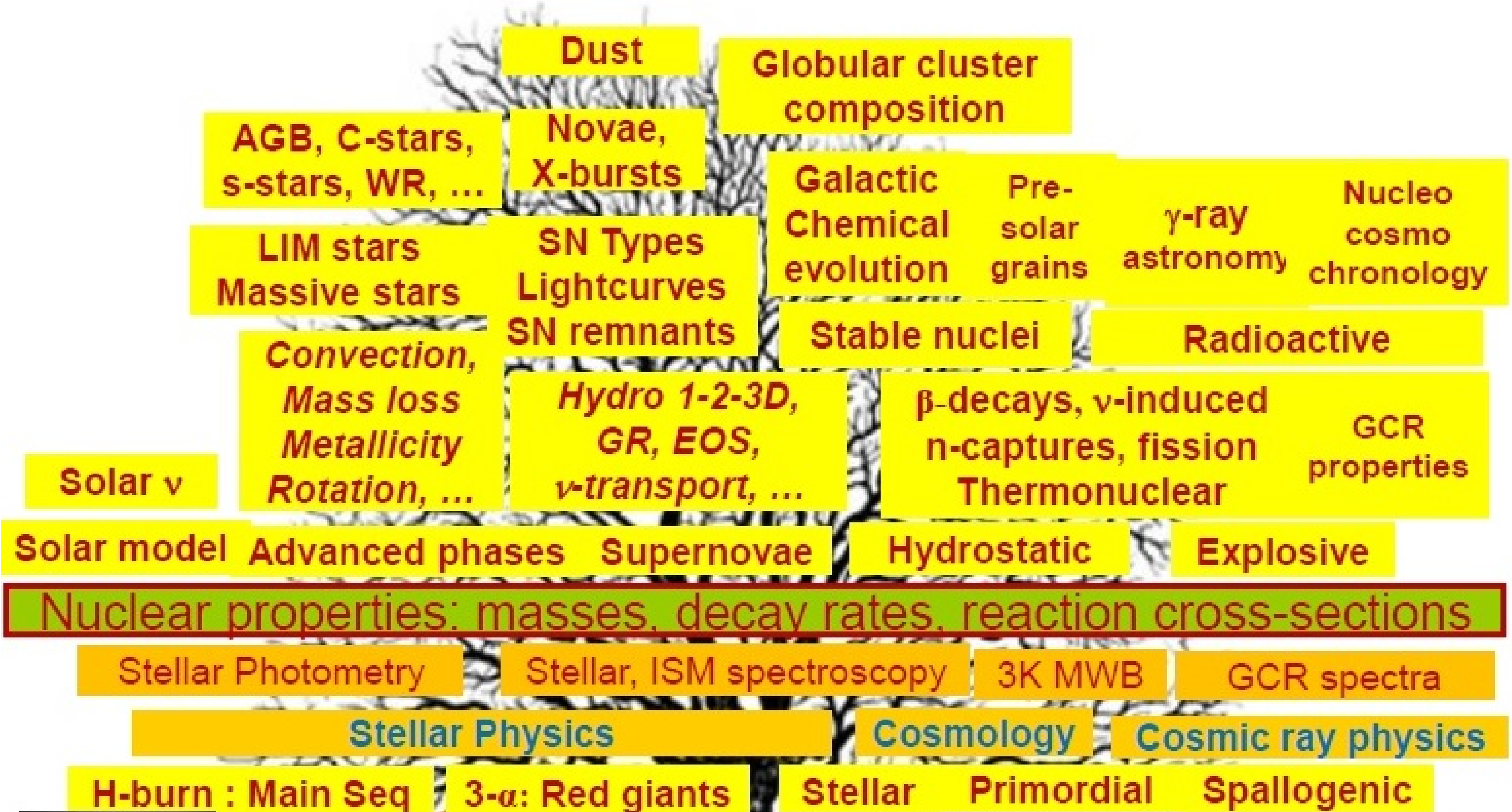
Neutron star

Black hole


abundances at solar system formation 4.5 Gy ago


William A. Fowler


Fred Hoyle



Dust

Globular cluster composition

AGB, C-stars, s-stars, WR, …

Novae, X-bursts

Galactic Chemical evolution

Pre-solar grains

γ-ray astronomy

Nucleo cosmo chronology

LIM stars Massive stars

SN Types Lightcurves SN remnants

Stable nuclei

Radioactive

Convection, Mass loss Metallicity Rotation, …

Hydro 1-2-3D, GR, EOS, ν-transport, …

β-decays, ν-induced n-captures, fission Thermonuclear

GCR properties

Solar ν

Solar model

Advanced phases

Supernovae

Hydrostatic

Explosive

Nuclear properties: masses, decay rates, reaction cross-sections

Stellar Photometry

Stellar, ISM spectroscopy

3K MWB

GCR spectra

Stellar Physics

Cosmology

Cosmic ray physics

H-burn : Main Seq

3-α: Red giants

Stellar

Primordial

Spallogenic

Stellar Energy

Origin of Elements

Nuclear Astrophysics

Nuclear Physics

Astrophysics